Probing the Physics of Star-Formation (ProPStar) \star

II. The first systematic search for streamers toward protostars

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ABSTRACT

Context. The detection of narrow channels of accretion toward protostellar disks, known as streamers, have increased in number in the last few years. However, it is unclear if streamers are a common feature around protostars that were previously missed, or if they are a rare phenomenon.

Aims. Our goals are to obtain the incidence of streamers toward a region of clustered star formation and to trace the origins of their gas, to determine if they originate within the filamentary structure of molecular clouds or from beyond.

Methods. We used combined observations of the nearby NGC 1333 star-forming region, carried out with the NOEMA interferometer and the IRAM 30m single dish. Our observations cover the area between the IRAS 4 and SVS 13 systems. We traced the chemically fresh gas within NGC 1333 with HC_3N molecular gas emission and the structure of the fibers in this region with N_2H^+ emission. We fit multiple velocity components in both maps and used clustering algorithms to recover velocity-coherent structures.

Results. We find streamer candidates toward 7 out of 16 young stellar objects within our field of view. This represents an incidence of approximately 40% of young stellar objects with streamer candidates when looking at a clustered star forming region. The incidence increases to about 60% when we considered only embedded protostars. All streamers are found in $HC₃N$ emission.

Conclusions. Given the different velocities between HC_3N and N_2H ⁺ emission, and the fact that, by construction, N_2H^+ traces the fiber structure, we suggest that the gas that forms the streamers comes from outside the fibers. This implies that streamers can connect cloud material that falls to the filaments with protostellar disk scales.

Key words. stars: circumstellar matter – stars: formation – stars: protostars – ISM: kinematics and dynamics

1. Introduction

Molecular clouds are composed of filaments [\(André et al. 2014;](#page-15-0) [Hacar et al. 2023\)](#page-16-0), which contain the cores where protostars and binary systems are born (e.g. [André et al. 2010;](#page-15-1) Off[ner et al.](#page-16-1) [2023\)](#page-16-1). These filaments are highly dynamic structures: gas has been observed to flow along them and they also accrete more gas from their surroundings (see [Hacar et al. 2023,](#page-16-0) for more details). Within filaments, molecular gas tends to be organized in velocity-coherent structures called fibers (e.g. [Hacar et al. 2017\)](#page-16-2). Observations of filaments and fibers highlight their relevance in directing molecular gas toward the sites of star formation.

Although astronomers have observed and modeled how mass flows within filaments and fibers in general, how that mass reaches the protostellar disk (on scales of ≤ 0.1 pc) is not well

understood. The way that mass reaches a protostellar system plays an important role in the star and planet formation process. For instance, infall from the envelope to the disk that is variable in time influences the accretion rate of the protostar, affecting its luminosity [\(Padoan et al. 2014;](#page-16-3) Kuff[meier et al. 2018\)](#page-16-4) and the chemical composition of the disk, represented, for instance, by the location of its snowline [\(Hsieh et al. 2019\)](#page-16-5) Numerical simulations also show that when infall from the envelope to the disk is heterogeneous in space, this can produce changes in disk structure, such as rings and gaps [\(Kuznetsova et al. 2022\)](#page-16-6) and second-generation disks misaligned with inner disks (Kuff[meier](#page-16-7) [et al. 2021\)](#page-16-7). Therefore, observing the mass flow from fiber scales (≥ 0.1 pc or ∼ 20 000 au) to disk scales (few ~ 100 au) is crucial to understand the influence of filaments and fibers in protostellar and disk properties.

In the last few years there has been a rise in the observations of streamers, defined as velocity-coherent, narrow structures that deposit their material to protostellar and protoplanetary disks [\(Pineda et al. 2023\)](#page-16-8). They are observed as asymmetries in the protostellar envelopes with total lengths between 500 au [\(Garufi](#page-16-9)

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[et al. 2022\)](#page-16-9) to even 10 000 au away from the protostar, beyond the natal protostellar core [\(Pineda et al. 2020\)](#page-16-10). These structures are different from fibers, which are velocity-coherent structures within the filaments. Streamers have been mostly detected and characterized toward embedded protostars (which are known as Class 0 and I sources, e.g. [Chou et al. 2016;](#page-16-11) [Valdivia-Mena et al.](#page-16-12) [2022,](#page-16-12) [2023;](#page-16-13) [Kido et al. 2023;](#page-16-14) [Aso et al. 2023\)](#page-15-2), but some streamers have been detected toward T Tauri sources (also known as Class II protostars) as well (e.g. [Ginski et al. 2021;](#page-16-15) [Garufi et al.](#page-16-9) [2022;](#page-16-9) [Harada et al. 2023\)](#page-16-16). Although mostly characterized toward low-mass young stellar objects (YSOs), recently accretion streamers have been discovered toward high-mass YSOs as well [\(Fernández-López et al. 2023\)](#page-16-17).

Although streamers appear increasingly common, their role in the larger puzzle of star formation is unclear.. One open questions is where do streamers come from? Numerical simulations show asymmetric accretion channels potentially generated by turbulence, both within the core itself (e.g. [Walch et al. 2010;](#page-16-18) [Seifried et al. 2013;](#page-16-19) [Hennebelle et al. 2020\)](#page-16-20), as well as coming from outside the natal core (Kuff[meier et al. 2017,](#page-16-21) [2023;](#page-16-22) [Heigl et al. 2024\)](#page-16-23). Observations of different species of molecular gas suggest that streamers can transport gas that comes from beyond the filaments themselves. For instance, [Valdivia-Mena](#page-16-13) [et al.](#page-16-13) [\(2023\)](#page-16-13) suggest that the observed gas falling toward fibers at scales of ∼ 20 000 au is connected to a streamer feeding disk scales in Barnard 5, but they cannot directly confirm this suggestion due to the different tracers and spatial resolutions of the observations.

Another point of debate is how frequent streamers actually are. The previously mentioned discoveries of streamers are mostly serendipitous. It is possible that accretion streamers are a common feature within protostellar envelopes of all ages, but they have not been characterized because the observations targeted other features, such as the disk or outflows (e.g. [Thieme](#page-16-24) [et al. 2022\)](#page-16-24). For instance, the the primary goal of the ALMA Large Program eDisk is to find substructures in young, embedded disks, yet their observations reveal streamers as well [\(Kido](#page-16-14) [et al. 2023;](#page-16-14) [Aso et al. 2023\)](#page-15-2). To this date, there has not been a systematic search for streamers toward YSOs. To better understand the role of streamers in our new picture of star formation, it is necessary to actively search for streamers in data that can also trace the larger fiber kinematics.

We present the first systematic study of gas flow from filament scales to individual protostellar envelopes with the explicit goal to search for streamers. This is part of the "Probing the Physics of Star-formation" (ProPStar) survey [\(Pineda et al.](#page-16-25) [2024\)](#page-16-25), where we explore the connection between the molecular gas within filaments and the circumstellar disk scales. We trace the flow of gas from a filament to YSOs, using a set of observations that allows us to characterize the flow of gas at fiber scales as well as toward individual protostars. The goal of this work is to search for streamers in a systematic fashion within a region where the kinematic properties of the fibers are known, to investigate the connection between fibers and streamers. For this purpose, we selected NGC 1333, one of the closest young embedded clusters, containing close to 150 YSOs [\(Gutermuth](#page-16-26) [et al. 2008\)](#page-16-26). It is the most active star forming region within the Perseus molecular cloud, located at 293 pc from Earth [\(Ortiz-](#page-16-27)[León et al. 2018;](#page-16-27) [Zucker et al. 2018\)](#page-16-28). Its high protostellar activity is reflected in the dozens of outflows stirring the local gas [\(Plunkett et al. 2013\)](#page-16-29). For this work, we select a region that covers two fibers with known kinematic properties which include the SVS 13 and the IRAS 4 protostellar systems. The observed area includes a total of 16 YSOs between Class 0 and Class II, enabling a first pilot study of incidence of streamers within an active star forming region. We refer to this whole area as the southeast (SE) filament.

This article is divided as follows: In Section [2,](#page-1-0) we describe the observations with the NOEMA and 30-m telescopes and their combination. In Section [3,](#page-2-0) we describe the methods used to decompose the observed line emission into velocity-coherent clusters for kinematic analyses. Section [4](#page-5-0) shows the resulting velocity structure of the fibers seen in the chosen gas tracers, the steps followed to find streamers in the data, and a description of the found candidates. We discuss our results and their physical interpretation in Section [5.](#page-14-0) Finally, in Section [6,](#page-15-3) we summarize the main results of our work.

2. Observations and Data reduction

We summarize the single dish and interferometric observations of NGC 1333 SE that we use from the ProPStar survey. The same observational setup is described in the ProPStar I paper [\(Pineda](#page-16-25) [et al. 2024\)](#page-16-25).

2.1. IRAM 30-m telescope

The single-dish observations were carried out using the 30-m telescope from IRAM, located in the Pico Veleta mountain in Spain. The observations are part of project 091-21 and were done during 2021 November 9, 10, and 11; and 2022 February 19, 20. We tuned the EMIR 090 receiver to cover the ranges between 72.3 – 78.8 GHz and 87.8 – 94.3 GHz. The FTS50 backend was employed, yielding a spectral resolution of 52.5 kHz. The complete map is $\approx 150'' \times 150''$ and was observed using an on-the-fly technique with position switching.

Data reduction was performed using the CLASS program of the GILDAS package^{[1](#page-1-1)}. The beam efficiency, B_{eff} , was obtained using the Ruze formula (available in CLASS), and was used to convert the observations into main beam temperatures, T_{mb} .

2.2. NOEMA

The observations were carried out with the IRAM NOrthern Extended Millimeter Array (NOEMA) interferometer as part of the the S21AD program using the Band 1 receiver, with the array in D configuration. The observations were taken on 2021 July 18, 19, and 21; August 10, 14, 15, 19, 22, and 29; and September 1. The mosaic consists on 96 pointings, centered at α_{J2000} =03^h29^m10.2s, δ_{J2000} =31°13′49.4″, which were observed
on four different scheduling blocks. We used the PolyFix coron four different scheduling blocks. We used the PolyFix correlator with a LO frequency of 82.505 GHz, an instantaneous bandwidth of 31 GHz spread over two sidebands (upper and lower) and two polarisations. The centers of the two 7.744 GHz wide sidebands are separated by 15.488 GHz. Each sideband is composed of two adjacent basebands of ∼3.9 GHz width (inner and outer basebands). In total, there are eight basebands which were fed into the correlator. The spectral resolution is 2 MHz throughout the 15.488 GHz effective bandwidth per polarization. Additionally, a total of 112 high-resolution chunks are placed, each with a width of 64 MHz and a fixed spectral resolution of 62.5 kHz. Both polarizations (H and P) are covered with the same spectral setup, and therefore the high-resolution chunks provide 66 dual polarisation spectral windows. In this work, we used the spectral windows containing the N_2H^+ *J* = 1 – 0 line and its hyperfine components (F_1F) at approximately 93 GHz,

¹ <http://www.iram.fr/IRAMFR/GILDAS>

Fig. 1: Peak temperature T_{peak} maps of the NOEMA and 30-m telescope observations. The black symbols represent YSOs in the region, summarised in Table [2:](#page-3-0) stars mark the positions of Class 0, I and 0/I protostars, circles represent Flat-spectra objects, and diamonds mark Class II sources. The protostars are labeled following Table [2.](#page-3-0) Left: N_2H^+ (1 – 0) $F_1F = 01 - 12T$ _{peak}. Right: HC₃N $(10 - 9) T_{peak}.$

and the HC₃N $J = 10 - 9$ line at approximately 91 GHz (Table [1\)](#page-3-1).

Data reduction was done using the CLIC program from the GILDAS package. We used the NOEMA pipeline to obtain the calibrated uv-tables, which we then combined with the singledish data as explained in the next section.

2.3. Combination of single-dish and interferometric data

The combination and imaging of the datasets previously described was done in mapping. The original 30-m data is resampled to match the spectral setup of the NOEMA observations. We use the task uvshort to generate the pseudo-visibilities from the 30-m data for each NOEMA pointing. The imaging is done with natural weighting, a support mask, and using the SDI deconvolution algorithm.

The final cubes are in K and have a channel resolution of approximately 0.2 km s⁻¹. The properties of the final data cubes are described in Table [1.](#page-3-1) The channel maps for each of the molecules are presented in Appendix [A.](#page-17-0) The observations cover an area of approximately $150'' \times 150''$. As we are interested in the kinematic structure of the filaments, we isolated one of the hyperfine structure emission lines of N_2H^+ *J* = 1 – 0, the $F_1F = 01 - 12$ component at 93176.2522 MHz. We chose this hyperfine component because it is the one that is the most separated from the other hyperfine structure lines, and therefore each individual peak can be interpreted as an individual velocity structure along the line of sight (Fig. 1 in [Caselli et al. 1995\)](#page-16-30). The N_2H^+ *J* = [1](#page-2-1) – 0, $F_1F = 01 - 12$ peak temperature map is in Fig. 1 left. The clean beam size for this molecule is approximately 5.8″ (1750 au). The mean rms of $N_2H^+(1_{01}-0_{12})$ cube is approximately 15.7 mJy beam⁻¹ (70 mK). For the rest of the article, we refer to the N_2H^+ *J* = 1 – 0, $F_1F = 01 - 12$ as N_2H^+ unless otherwise stated.

The HC₃N $J = 10 - 9$ peak temperature map is shown in Fig. [1](#page-2-1) right. The clean beam size at this frequency is approximately 4.9''(1400 au). The mean rms of the HC₃N $J = 10 - 9$

emission cube is approximately 12 mJy beam^{-1} (80 mK) and increases at the edges of the map due to the primary beam response. We refer to $HC_3N (10-9)$ line emission as HC_3N in the rest of this work, as is the only HC_3N emission line presented.

The YSOs that are found within the observed area and their general characteristics are summarised in Table [2.](#page-3-0) We plot these protostars and Class II sources in Fig. [1](#page-2-1) and label each one so that the reader has a quick reference of where each one is in for the following analyses. We include IRAS 2A even though our map does not cover the protostar as it has a strong outflow that needs to be taken into account to understand the gas kinematics of the region.

3. Line decomposition methods

The peak temperature maps shown in Fig. [1](#page-2-1) do not fully reflect the complexity of the obtained data. Both molecular lines show several emission peaks, some located at or within a beam of protostars, others apparently unrelated to protostellar sources. However, an initial inspection of the data cube shows that there are several locations in the map where the spectra have two or three velocity components along the line of sight. Previous molecular gas observations done both with similar resolution to our data [\(Dhabal et al. 2019\)](#page-16-31) as well as with single-dish telescopes [\(Hacar](#page-16-2) [et al. 2017;](#page-16-2) [Dhabal et al. 2018;](#page-16-32) [Chen et al. 2020b\)](#page-16-33) show that there are multiple velocity components in this region. In particular, [Hacar et al.](#page-16-2) [\(2017\)](#page-16-2) and [Chen et al.](#page-16-33) [\(2020b\)](#page-16-33) find that these velocity components are due to three velocity-coherent fibers of gas in the region, two of which run in the northwest to southeast direction that overlap in line of sight, although the exact structure identification is different for both works (we discuss the differences further in Appendix [C\)](#page-17-1). For the kinematic analysis of the gas surrounding the protostars, it is necessary to separate these velocity components.

In this section, we describe the procedure used to separate the velocity components of NGC 1333 SE. In summary, we first separated the individual velocity components of each spectrum

Table 1: Spectral lines observed in the high-resolution chunks used in this work

N_2H^+ 6.1×10^{4} 70 $J = 1 - 0$, $F_1 F = 01 - 12$ 93176.2522 4.47 5.76×5.48 1.6×10^{5} 90979.0230 24.01 80 HC ₃ N 4.89×4.69 $J = 10 - 9$	Molecule	Transition	ν (MHz)	E_{up} (K)	n_{crit} (cm ⁻³)	beam size	rms(mK)

Notes. Frequencies taken from the CDMS catalog. Critical densities taken from [Shirley](#page-16-34) [\(2015\)](#page-16-34). The rms is taken from the combined NOEMA and 30-m data cubes.

Table 2: Properties of the protostellar objects found within the N_2H^+ and HC_3N maps.

Source	RA	DEC	Other Names	Class ^a	Multiple? $\frac{b}{b}$	$v_{LSR}^{\ c}$	Outflow P.A. d	Disk i^e
	(J2000)	(J2000)				$(km s^{-1})$	$(^\circ)$	(°)
IRAS4A	03:29:10.54	$+31:13:30.93$	Per-emb 12	θ	Y	6.9 ± 0.005^1	19 ⁶	35
IRAS4B	03:29:12.02	$+31:13:08.03$	Per-emb 13	θ	Y	7.1 ± 0.009 ¹	176 ⁶	49
Per-emb-14	03:29:13.55	$+31:13:58.15$	NGC 1333 IRAS4C	θ	N	7.9 ± 0.03^1	96^{7}	64
Per-emb-15	03:29:04.06	$+31:14:46.24$	RNO15-FIR, SK 14	θ	N	6.8 ± 0.01^1	-35^{5}	
Per-emb-27	03:28:55.57	$+31:14:37.03$	NGC 1333 IRAS2A	0/I	Y	8.1 ± 0.02^1	$105^{5,f}$	
Per-emb-36	03:28:57.37	$+31:14:15.77$	NGC 1333 IRAS2B	L	Y	6.9 ± 0.02^1	24^{5}	
Per-emb-44	03:29:03.76	$+31:16:03.81$	SVS ₁₃ A	0/I	Y	8.7 ± 0.02 ¹	140^{8}	
SVS ₁₃ C	03:29:01.97	$+31:15:38.05$		Ω	Y	8.9 ± 0.02 ¹	8 ⁸	75
SVS ₁₃ B	03:29:03.08	$+31:15:51.74$		$\mathbf{0}$	Y	8.5 ± 0.01^1	170 ⁶	61
IRAS4B2	03:29:12.84	$+31:13:06.89$	NGC 1333 IRAS4B'	Ω	Y		-99^{9}	
EDJ2009-183	03:28:59.30	$+31:15:48.41$	ASR 106	Flat	Y	8.69 ²		
EDJ2009-173	03:28:56.96	$+31:16:22.20$	ASR 118, SVS15	\mathbf{I}	N	9.08 ²		
VLA ₃	03:29:03.00	$+31:16:02.00$		Ω				
SK 15	03:29:06.50	$+31:15:38.60$	ASR 6, HRF 50	L		7.95^{3}		
ASR ₃	03:29:02.16	$+31:16:11.40$	[GMM2008] 76	\mathbf{I}				
ASR 53	03:29:02.16	$+31:16:11.40$	[GMM2008] 89	\mathbf{I}		13.5 ± 3.4^4		
ASR 54	03:29:09.41	$+31:14:14.10$	[GMM2008] 135	\mathbf{I}				

Notes. We only describe the properties of protostars that are on regions where there is HC_3N emission or the outflow lobe is within the HC_3N emission map. (*a*) SED classification from the literature [\(Enoch et al. 2009a;](#page-16-35) [Evans et al. 2009\)](#page-16-36). We indicate also when a source is a Class 0/I according to the classification in the VANDAM sample [\(Tobin et al. 2018\)](#page-16-37). ^(b) Y if the source contains more than one protostar, N if it is single [\(Tobin et al. 2018\)](#page-16-37) (*c*) Obtained from available molecular tracer observations of cores or from spectroscopic surveys. References: (1) [Stephens](#page-16-38) [et al.](#page-16-38) [\(2019\)](#page-16-38), (2) [Foster et al.](#page-16-39) [\(2015\)](#page-16-39), (3) [Imai et al.](#page-16-40) [\(2018\)](#page-16-40), (4) [Kounkel et al.](#page-16-41) [\(2019\)](#page-16-41). (*d*) Angle of the red lobe from North to East. References: (5) [Stephens et al.](#page-16-42) [\(2017\)](#page-16-42), (6) [Lee et al.](#page-16-43) [\(2016\)](#page-16-43), (7) [Zhang et al.](#page-16-44) [\(2018\)](#page-16-44), (8) [Plunkett et al.](#page-16-29) [\(2013\)](#page-16-29), (9) [Podio et al.](#page-16-45) [\(2021\)](#page-16-45) (*e*) Disk inclination with respect to the plane of the sky from [Segura-Cox et al.](#page-16-46) [\(2018\)](#page-16-46). ^(*f*) In IRAS 2A, there are two outflows perpendicular to each other. Here we list the most collimated one, where the red lobe is included in the footprint of our $HC₃N$ data.

within the N_2H^+ and HC_3N maps by fitting multiple Gaussians. Next, we clustered the velocity components to recover individual, velocity-coherent fibers in the region. This process allowed us to analyze the velocity structure around the protostars more easily.

3.1. Identification of velocity components

Fitting a model using a simple minimization algorithm, such as least squares, has issues in identifying the best parameters for faint signals. Moreover, quantifying how much better a more complex model (i.e. with more parameters) is compared to a simpler one is a delicate task (e.g. [Protassov et al. 2002\)](#page-16-47). A quick look at the data cubes indicates that there are lines of sight where the possible second component has a signal-to-noise ratio (S/N) ∼ 5. Therefore, a simple least squares fit followed by a statistic evaluation might not be enough to effectively separate the kinematic components within the filaments.

Bayesian model selection using nested sampling can solve these problems. Nested sampling is a parameter exploration algorithm that evaluates how probable the combination of parameters is, therefore allowing parameter estimation and model comparison, even in low S/N data [\(Skilling 2004,](#page-16-48) [2006;](#page-16-49) [Sokolov](#page-16-50) [et al. 2020\)](#page-16-50). This algorithm returns the Bayesian Evidence Z of a model, which is a likelihood integral of the parameter values:

$$
\mathcal{Z} = \int p(\theta) \mathcal{L}(\theta) d\theta,\tag{1}
$$

where θ are the probable parameter values, $p(\theta)$ is the probability density function, and $\mathcal{L}(\theta)$ is the likelihood function. From \mathcal{Z} , we can then calculate the relative probability K_{n-1}^n that a number of components *n* returns a better fit than *n* − 1 components:

$$
K_{n-1}^{n} = \frac{P(M_n)\mathcal{Z}_n}{P(M_{n-1})\mathcal{Z}_{n-1}} = \frac{\mathcal{Z}_n}{\mathcal{Z}_{n-1}},
$$
\n(2)

where $P(M_n)$ is the probability *a priori* of a model with *n* Gaussian components (for our case). We assume that all competing models are equally likely *a priori*, so $P(\mathcal{M}_n) = P(\mathcal{M}_{n-1})$.

We used the gaussian model in the Python package pyspeckit to fit the spectra with 1 and 2 Gaussian components for N_2H^+ and up to 3 Gaussian components for HC_3N . The third component in the case of HC_3N is to account for out-flow emission. We used pymultinest [\(Buchner et al. 2014\)](#page-16-51), a package designed to use the Fortran-based package MultiNest in Python [\(Feroz et al. 2009;](#page-16-52) [Feroz & Hobson 2008\)](#page-16-53), to run the

Fig. 2: Resulting number of individual velocity components fitted along each line of sight for N_2H^+ (left) and HC_3N (right). The red symbols represent protostellar objects in the region as in Fig. [1.](#page-2-1) The gray regions correspond to pixels with SNR≤ 2 or areas without data. The black semitransparent lines show the spines of the fibers in $NH₃$ [\(Chen et al. 2020b\)](#page-16-33). The red ellipse at the bottom left corner of each plot represents the beam size.

nested sampling. In conjunction to pymultinest, we used the Python package pyspecnest^{[2](#page-4-0)} [\(Sokolov et al. 2020\)](#page-16-50) to wrap the pyspeckit fitting classes so as to use them in pymultinest. We followed the steps and code used in [Sokolov et al.](#page-16-50) $(2020)^3$ $(2020)^3$ $(2020)^3$.

From this process, we obtained the best fit parameters and their uncertainties for every spectra in the cubes with SNR> 2. We adopted a conventional cut for *K* of $\ln K_{n-1}^n > 5$, so if $\ln K^n \le 5$ we select $n-1$ components. We use this criterion ln K_{n-1}^n ≤ 5, we select *n* − 1 components. We use this criterion to obtain a map with the number of components *n* in each pixel. There are small regions (consisting of less than 100 pixels) of emission in this map that are surrounded by noisy data, so we ran two additional steps to reduce the variation of *n* due to noise.

First, we eliminated small regions of consisting of *n* components surrounded by *n* − 1 components or emission surrounded by noise $(n = 0)$, which we call islands, by replacing them with their resulting $n - 1$ fit results. Islands of less than 100, 7 and 2 pixels for the cases of $n = 1$, $n = 2$ and $n = 3$ respectively were replaced with the results from $n - 1$ fit. These thresholds are based on a visual inspection of the resulting fits, where islands with a size smaller than the threshold tend to fit noise peaks outside the observed emission range (5 to 10 km s[−]¹). Nevertheless, the exact size of the islands does not influence the final results.

Secondly, we filtered the fit results using the parameter uncertainties. For our analysis, we only used the fit results that have an uncertainty in central velocity smaller than the channel size. We evaluated each individual fit component to determine its quality. If the central velocity of one component has a larger uncertainty than the channel width, that component was eliminated, without affecting the other components that might be present in the same spectrum. The *n* map is updated accordingly,

for instance, if $n = 3$ in a pixel and one component was eliminated, we replaced $n = 2$ in that position. After this filtering, we obtained central velocities that have uncertainties between 0.02 and 0.03 km s⁻¹ for N₂H⁺, and between 0.03 and 0.05 km s⁻¹ in for HC_3N .

Figure [2](#page-4-2) shows the final number of Gaussian components per spectrum in the N_2H^+ (left) and HC_3N (right) emission cubes. We required up to 2 components for N_2H^+ and up to 3 in HC_3N to adequately fit the spectra, where the third Gaussian component in all cases was used to fit extended wings of emission in the HC_3N spectra. Section [4.1.2](#page-5-1) describes these wings in more detail.

3.2. Clustering of velocity structures

We clustered the individual velocity components of each of the molecular gas emissions to compare HC_3N emission with N_2H^+ emission. This comparison is not straightforward due to the presence of multiple velocity components distributed differently in both molecules. Figure [2](#page-4-2) left shows that N_2H^+ has mostly two velocity components at the center of the map, located similarly to the overlapping fibers found with lower resolution data $NH₃$ (approximately 30'') by [Chen et al.](#page-16-33) [\(2020b\)](#page-16-33). However, HC_3N does not show this overlap, with regions of two and three velocity components scattered throughout the map (Fig. [2](#page-4-2) right). Therefore, it is not clear which Gaussian component in HC₂N should be compared with what N_2H^+ component. By clustering the Gaussian peaks, the comparison between both molecules is possible.

We clustered the components using a density-based clustering algorithm because it is not possible to separate them manually using a simple velocity threshold. If we grouped all velocity components larger than a certain v_{LSR} together, it is possible we

² https://github.com/vlas-sokolov/pyspecnest

³ https://github.com/vlas-sokolov/bayesian-ngc1333

Fig. 3: Results of the clustering algorithms for the Gaussian components of N_2H^+ and HC_3N . The black circles at the bottom left corners represent the beam size of the respective data. The scalebar at the top left corner represents a length of 10 000 au. The black semitransparent lines show the spines of the fibers in NH₃ [\(Chen et al. 2020b\)](#page-16-33). Left: cluster groups for N_2H^+ , labeled as red and blue representing the more redshifted and blueshifted groups, respectively. Right: cluster groups for HC_3N , labeled 0 to 6.

separated emission that is connected with the grouped emission but that has a large velocity gradient, and so has components with a lower v_{LSR} than the threshold. A look at the velocities obtained after the nested sampling (Fig. [B.1](#page-19-0) and [B.2\)](#page-19-1) shows that both molecules present complex velocity structures.

We describe all the details of the clustering process for each molecule in Appendix [B.](#page-17-2) In summary, we used Hierarchical Density-Based Spatial Clustering of Applications with Noise (HDBSCAN) to find clusters of emission in position-positionvelocity space. We excluded from the clustering the Gaussian components with $\sigma_v > 1 \text{ km s}^{-1}$, as these trace outflow cones
in HC. N (Sect 4.1.2) Figure 3 shows the resulting clusters and in $HC₃N$ (Sect. [4.1.2\)](#page-5-1). Figure [3](#page-5-2) shows the resulting clusters and their labels. The complete set of properties (peak temperature, velocity and velocity dispersion) are in Fig. [B.2](#page-19-1) for HC_3N and in Fig. [B.1](#page-19-0) for N_2H^+ clusters.

4. Results and analysis

4.1. Properties of the NGC 1333 SE gas

4.1.1. Velocity structure of N_2H^+ and HC_3N

We recovered the fiber structure observed in previous works with HDBSACAN (Sect. [3.2\)](#page-4-3) in N_2H^+ emission. HDBSCAN recovers two clusters: we named cluster 0 the blueshifted (blue) fiber and cluster 1, the redshifted (red) fiber as these resemble the fibers found in [Chen et al.](#page-16-33) [\(2020b\)](#page-16-33), where one has a larger v_{LSR} than the other. The distribution of the clusters is more similar to the fibers found in NH_3 by [Chen et al.](#page-16-33) [\(2020b\)](#page-16-33) than for previous N_2H^+ decomposition by [Hacar et al.](#page-16-2) [\(2017\)](#page-16-2). We discuss the differences between the structure characterization in our work with respect to [\(Hacar et al. 2017\)](#page-16-2) in Appendix [C.](#page-17-1) Figure [4](#page-6-0) top shows the v_{LSR} of each of the fibers. The redshifted fiber has

an average < v_{LSR} >= 7.87 km s⁻¹ and the blueshifted fiber,
< v_{LSR} >= 6.95 km s⁻¹. The red cluster captures additional $\langle v_{LSR} \rangle = 6.95$ km s⁻¹. The red cluster captures additional
emission toward the northeast of the map part of the Northeast emission toward the northeast of the map, part of the Northeast (NE) filament (as named by [Dhabal et al. 2019\)](#page-16-31), and toward the west, which is part of the extended emission not covered by our observations but observed in N_2H^+ emission [\(Hacar et al. 2017\)](#page-16-2).

The clustering in HC₃N, unlike N_2H^+ , separates the emission into 7 clusters. There are no HDBSCAN parameters that give a similar clustering to N_2H^+ . We grouped the HC_3N clusters into redshifted and blueshifted groups according to their velocity with respect to the average velocity of the N_2H^+ clusters, so as to compare their velocities (Sect. [4.1.3\)](#page-7-0). Clusters in HC_3N which have an average velocity closest to the redshifted N_2H^+ fiber are altogether called redshifted fiber, and clusters closer in velocity to the blue N_2H^+ are grouped under the blueshifted fiber. The vLSR of the resulting groups are shown in Fig. [4.](#page-6-0)

4.1.2. Correlation of HC_3N emission with outflows

Figure [5](#page-7-1) left shows the integrated intensity of $HC₃N$ together with the ¹²CO $J = 2 - 1$ integrated emission maps toward the protostars, available from the MASSES survey [\(Stephens et al.](#page-16-38) [2019\)](#page-16-38). We show the spectra at the locations where there are three Gaussians in the HC_3N cube (toward the north and south of IRAS 4A and toward the east of Per-emb 36) in Fig. [5](#page-7-1) right. The spectra at these locations present extended wings: at the south of IRAS 4A, the wing is blueshifted with respect to the emission peak, and toward the north of IRAS 4A and the east of IRAS 2B, the wing is redshifted. All of these wings were fitted with one or two broad ($\sigma_{\rm v} > 2 \text{ km s}^{-1}$) Gaussian components. The pres-
ence of these wings and their velocity (blueshifted or redshifted) ence of these wings and their velocity (blueshifted or redshifted) coincide with the outflows shown in ¹²CO $J = 2 - 1$.

Fig. 4: Resulting velocity groups for N_2H^+ and HC_3N emission after the clustering process. The left plots show the v_{LSR} of the redshifted fiber, whereas the right plots show the same for the blueshifted fiber. The blue contour marks the area occupied by the blueshifted fiber and the red contour, the redshifted fiber. The black symbols represent protostellar objects in the region as in Fig. [1.](#page-2-1) The black ellipse at the bottom-left corner of both plots represents the beam size. The black lines represent the velocity gradient directions measured at each position, with their size representing the gradient magnitude. The red circle at the top right corner of the left plot shows the size of the sampled area to calculate the gradients. The panels do not have the same spatial scales. Top: N_2H^+ emission v_{LSR} for each of the clusters. Bottom: HC₃N emission v_{LSR} after grouping the clusters according to their velocities with respect to the N_2H^+ clusters.

Fig. 5: Correlation between the HC₃N emission and outflows. Left: Integrated intensity map of HC₃N (10 – 9) between 5 and 10 km s⁻¹. The red symbols represent protostellar objects in the region as in Fig. [1.](#page-2-1) The red and blue contours correspond to the ¹²CO red and blue outflow lobes, respectively, obtained from the MASSES survey [\(Stephens et al. 2019\)](#page-16-38). The labels indicate where each of the spectra is taken from. Right: HC₃N (10 – 9) spectra at the locations of the IRAS 4A outflow and at the bright emission to the east of IRAS 2B. For each location, we take the spectrum at an individual pixel. The blue, green and red curves represent the individual three Gaussians fitted at each position, and the red dashed curve represents the sum of all the Gaussians.

The wings in the spectra close to IRAS 4A are caused by the outflow driven by IRAS 4A, but the bright $HC₃N$ redshifted wing to the east of IRAS 2B coincides with the tip of the outflow lobe of IRAS 2A (also known as Per-emb 27). IRAS 2A is a binary system that has two outflows, one collimated in the east-west direction and a wider one in the north-south orientation [\(Plunkett et al. 2013\)](#page-16-29). The redshifted lobe of the collimated outflow is toward the east and shows intense ${}^{12}CO$ emission at the same location of the brightest peak of the $HC₃N$ integrated intensity map. Part of this redshifted ¹²CO emission is at the eastern edge of the IRAS 2B MASSES image and the other part is at the western edge of the Per-emb 15 MASSES map. As this emission does not coincide with other protostars or peaks in N_2H^+ emission, we conclude this HC_3N enhancement is a bow shock due to the outflow from IRAS 2A impacting the cloud material. In the HC_3N channel maps, between 6.5 and 8.8 km s⁻¹ (Fig. [A.1\)](#page-18-0), the images show a v shape at the position marked in Fig. [5](#page-7-1) as IRAS 2B east. We suggest this is due to the collision of the outflow lobe with cloud material, which produces a bow shock in the gas and enhances the presence of HC_3N .

Not all of the outflows cause wings in HC_3N emission. HC_3N is enhanced approximately along the IRAS 4C and IRAS 4B2 outflows (Fig. [5\)](#page-7-1), but the directions of the enhancements are not over the position angles of the corresponding outflows. In the SVS 13 system, HC_3N emission is brighter at the locations of the protostars and extends all along the ridge that joins the protostars together, but does not present extended wings in the

spectra. Moreover, comparing the ${}^{12}CO$ outflows from MASSES with the integrated HC_3N emission (without the outflow wings) in Fig. [5,](#page-7-1) HC_3N seems to be brighter *outside* of the regions covered by ¹²CO, with the exception of the IRAS 2A and IRAS 4A outflows. Therefore, the outflows influence HC_3N emission differently depending on the location, but in general, HC_3N seems to be enhanced at the outflow cavities, while tracing extended material in general.

4.1.3. Comparison between HC_3N and N_2H^+ velocities

We compared the extent and v_{LSR} of both molecules independently for each fiber. N_2H^+ is used as a proxy for the dense gas, as in Perseus N_2H^+ correlates well with locations of dense cores, where CO is depleted [\(Johnstone et al. 2010\)](#page-16-54). Previous works in this region show that N_2H^+ follows the same structures as dust and NH_3 , which means that N_2H^+ traces of the physical structure of each fiber (Hacar et al. 2017 ; [Dhabal et al. 2019\)](#page-16-31). HC₂N is a known "early type" molecule together with other carbonchain molecules (such as CCS and cyclic- C_3H_2), meaning that in chemical models of molecular cloud collapse, these molecules appear at early times in comparison to, for example, nitrogenbearing molecules [\(Suzuki et al. 1992;](#page-16-55) [Bergin & Tafalla 2007\)](#page-15-4). We used the channel maps together with the spectral decomposition made in Sect. [3](#page-2-0) to describe the gas structure.

In general, HC_3N structure consists of several small peaks in emission connected via less bright, extended emission. HC_3N does not cover the full extent of the NGC 1333 SE filament nor each of the N_2H^+ fibers: the area covered by HC_3N emission is smaller than the one covered by N_2H^+ emission, as seen in Fig. [2,](#page-4-2) and is particularly scattered toward the filament tail in the south. HC_3N is detected inside the area defined for both fibers with N_2H^{\dagger} emission: in the redshifted fiber, HC₃N is roughly detected along the center, whereas in the blueshifted fiber $HC₂N$ is detected only toward the west side of the fiber.

We compared the velocities traced by both molecules in each fiber. Figure [6](#page-9-0) left shows the velocity difference between HC_3N and N_2H^+ $\delta v_{LSR} = v_{LSR,HC_3N} - v_{LSR,N_2H^+}$. We reprojected the HC₃N images of each velocity component to the pixel grid of the N_2H^+ cube, because the pixel size and beam is larger. Most of the redshifted fiber has HC_3N velocities that are larger than the velocity of N_2H^+ gas. The exception to this is a band that runs east to west below the SVS 13 area. The blueshifted fiber shows considerable differences between the molecules' velocities close to IRAS 4A and 4B, but in the rest of the gas, there is no apparent difference.

We evaluated if the values of δv_{LSR} are significant using the Kernel Density Estimate (KDE) of the differences between the vLSR of both molecules. This KDE is in Fig. [6](#page-9-0) right. The median difference is 0.11 km s⁻¹ for the redshifted fiber and 0.01 km s⁻¹ for the blueshifted fiber. δv_{LSR} for the red fiber is larger than the sum of the uncertainties in both $\rm v_{LSR,HC_3N}$ and $\rm v_{LSR,N_2H^+}$ obtained from the Gaussian fitting (0.08 km s^{-1}) , and therefore show that the gas in both molecules presents different velocities. This is not the case for the blue fiber.

4.1.4. Velocity gradients within the fibers

We calculated the gradients present in the N_2H^+ and HC_3N velocity fields. The local velocity gradients describe the fluctuations of gas motion within the fibers, and are of assistance to identify infall toward individual YSOs or binaries. We used a similar approach as used in [Chen et al.](#page-16-56) [\(2022\)](#page-16-56) and [Valdivia-](#page-16-13)[Mena et al.](#page-16-13) [\(2023\)](#page-16-13). We calculated the local velocity gradient by fitting a plane centered at one pixel, with a width of 2 beams, so that we capture gradients in uncorrelated pixels. The plane is defined as:

$$
v_{LSR} = Ax + By + v_c,
$$
\n(3)

where *A* is the slope in the x (right ascension) direction and *B* in the y (declination) direction. We used the velocity gradient implementation within the velocity_tools package^{[4](#page-8-0)}.

The resulting gradient orientations and magnitudes are plot-ted over the v_{LSR} of each fiber in Fig. [4](#page-6-0) for HC_3N and N_2H^+ . In both fibers, the velocity gradient field for HC_3N is more randomly aligned than for N_2H^+ . This indicates that the gas traced by $HC₃N$ presents larger changes at local scales, in part due to the effect of outflows in the region.

Both N_2H^+ and HC_3N show large (about 20 km s⁻¹ pc⁻¹) gradients within the fibers, almost completely perpendicular to the filament, in regions where there are no protostars. In particular, the velocity gradients toward the south of the redshifted fiber in both molecules seem well ordered. Between IRAS 4 and SVS 13, the velocity gradients between both molecules are different. The N_2H^+ gradient vectors are mostly perpendicular to the length of the filament in the blueshifted fiber, but with a lower magnitude ($\leq 10 \text{ km s}^{-1} \text{ pc}^{-1}$). In the redshifted fiber, there is a region with a sudden change in velocity that produces local, perpendicular gradients of about 20 km s⁻¹ pc⁻¹, but in the rest

of the region the local gradients are smaller than 5 km s−¹ pc−¹ and do not show clear patterns. On the other hand, $HC₃N$ emission presents strong variations in the redshifted fiber, up to 30 km s−¹ pc−¹ , but apparently randomly oriented, and roughly perpendicular with magnitudes between 10 and 20 km s⁻¹ pc⁻¹ in the blueshifted fiber.

4.2. Streamer candidates

We analyzed the molecular tracers surrounding the YSOs in search for signatures of streamers. The observations in this work are designed to follow the flow from the larger scales of the filaments to the smaller, protostellar scales. We list the general properties of the YSOs within our field of view in Table [2.](#page-3-0)

The search for streamers in the region was done using the following signatures: we first looked for velocity gradients in our HC_3N and N_2H^+ maps where the difference between the vLSR of our observations and the protostar increases as the distance between the gas and the protostar decreases. This is a common characteristic of streamers observed towards protostars and pre-main sequence stars (e.g. [Thieme et al. 2022;](#page-16-24) [Valdivia-Mena](#page-16-12) [et al. 2022,](#page-16-12) [2023;](#page-16-13) [Hsieh et al. 2023b\)](#page-16-57). We analyzed both the maps of the velocity gradients within clusters (Fig. [4\)](#page-6-0) as well as the central velocities fitted to regions that were not clustered in Sect. [3.2.](#page-4-3) For embedded protostars (Class 0, 0/I and I), we used the vLSR reported in the MASSES survey [\(Stephens et al. 2019\)](#page-16-38) as protostellar velocity, obtained using a Gaussian fit to SMA $C^{18}O$ (2-1) emission observations within an area of 1.2′′. If the velocity was not available in MASSES, we used the v_{LSR} reported in single-dish observations of DCN [\(Imai et al. 2018\)](#page-16-40). For Class II sources, we used reported v_{LSR} values from APOGEE spectra [\(Foster et al. 2015;](#page-16-39) [Kounkel et al. 2019\)](#page-16-41). When the velocity of the protostars not available from the previous observations, we adopted the velocity of N_2H^+ at the source location. Then, we determined if the velocity gradient comes from a preferential direction and is not a radial velocity gradient centered at the protostar. Finally, we observed if there is any elongated structure, such as bright lanes or peaks not centered at the protostar's location, in T_{MB} . We defined the region that makes up the potential streamer manually, while including emission within the S/N $= 10$ contours.

We constructed sub-cubes of HC_3N and N_2H^+ emission and images of the Gaussian component properties (Sect. [3.1\)](#page-3-2) and clusters (Sect. [3.2\)](#page-4-3), centered at each protostar and including all pixels within a 10, 000 au radius. We chose 10, 000 au based on the longest confirmed streamer found to date [\(Pineda et al.](#page-16-10) [2020\)](#page-16-10), which is characterized using HC_3N emission. We excluded from our analysis the broad components correlated with outflow activity seen in Sect[.4.1.2.](#page-5-1) We expected HC_3N to reveal streamer motion, as streamers have been mostly observed in carbon-bearing molecules [\(Pineda et al. 2020\)](#page-16-10), and N_2H^+ can give information about the surrounding envelope, but we also checked the N_2H^+ velocity gradients. The beam of our HC_3N data is 4.9′′, which corresponds to a length of almost 1500 au. Based on this, we are able to determine the presence of streamers (or asymmetric infall) using HC₃N that are at least ~ 1500 au in projected length. We note that streamers can be longer in reality, but we are limited by their projected length in the plane of the sky unless we are able to model their infall kinematics, which was evaluated in each individual case. Therefore, our HC_3N data has the potential to reveal asymmetric infall of sizes typical of previously discovered streamers toward Class 0 and I protostars [\(Thieme et al. 2022;](#page-16-24) [Valdivia-Mena et al. 2022,](#page-16-12) [2023\)](#page-16-13). There is no detected HC_3N emission around ASR 53, EDJ2009-183

⁴ https://github.com/jpinedaf/velocity_tools

Fig. 6: Difference between the velocity components of each fiber in NGC 1333 SE, as recognized after the clustering. Left: Resulting difference after subtracting the N₂H⁺ v_{LSR} from the HC₃N v_{LSR} of the clusters belonging to the redshifted fiber. Middle: Resulting difference after subtracting the \tilde{N}_2H^+ v_{LSR} from the $\tilde{HC_3N}$ v_{LSR} of the clusters belonging to the blueshifted fiber. Note that the blueshifted fiber map is zoomed into the region covered by the emission. Right: two-dimensional KDE of the HC₃N v_{LSR} versus N_2H^+ v_{LSR} . The dotted line represents the location where the velocities are equal, and the dash-dotted line, where the HC_3N v_{LSR} is larger by 0.11 km s^{-1} (the median difference for the redshifted fiber).

(ASR 106) and EDJ2009-173 (ASR 118), and Per-emb 27 (IRAS 2A) is out of the imaged area, so these sources are not analyzed.

We found streamer candidates toward 7 out of the 16 YSOs from Table [2.](#page-3-0) In the following, we present the analysis of each individual protostar that has a streamer candidate. We present IRAS 4A first as it is the only one where there is enough resolution and available information about its stellar properties to confirm its infall motion using a free-fall model. Figure [7](#page-10-0) shows the peak brightness and velocity of HC_3N and N_2H^+ , together with a streamer trajectory solution zoomed at the selected 10, 000 au scale. Figure [8](#page-11-0) shows the same HC_3N properties (except for the free-fall model) for the rest of the streamer candidates, and Fig. [D.1](#page-20-0) shows their N_2H^+ observed properties. We found no evident streamer candidates in IRAS 4B2, IRAS 4C (Per-emb 14), SVS 13B, SVS 13C and ASR 3.

4.2.1. IRAS 4A

IRAS 4A is a Class 0 protostar that drives a strong outflow, described in several previous works (e.g., [Plunkett et al. 2013\)](#page-16-29). We also observed the effect of this outflow in the HC_3N emission (Sect. [4.1.2\)](#page-5-1). We examined the different cutout images for the surroundings of IRAS 4A and found a region with a velocity gradient consistent with infall motion. There is an increase in T_{MB} toward the west of the protostar in the HC_3N cluster no. 6, part of the blueshifted fiber. We show the T_{MB} of this cluster in Fig. [7](#page-10-0) top left. This region has a velocity gradient from west to east, which we show in Fig. [7](#page-10-0) bottom left: at approximately 2500 au from IRAS 4A, HC₃N gas has a v_{LSR} = 7.0 km s^{−1},
whereas closer to the protostar it reaches $y_{\text{c}} = 6.85$ km s^{−1} whereas closer to the protostar it reaches $v_{LSR} = 6.85 \text{ km s}^{-1}$.
The uncertainty in the fitted velocity is 0.03 km s^{−1} on average The uncertainty in the fitted velocity is 0.03 km s^{-1} on average for this region, so the difference between both extremes is significant. Therefore, we defined this region as a streamer candidate in $HC₃N$ emission.

We were able to confirm that this velocity profile is consistent with free-fall infalling gas toward the protostar. Using the methods detailed in [Pineda et al.](#page-16-10) [\(2020\)](#page-16-10) and [Valdivia-Mena et al.](#page-16-12) [\(2022\)](#page-16-12), we modeled the velocity gradient in the image plane and the velocity plane, assuming a free-fall, ballistic trajectory from [Mendoza et al.](#page-16-58) [\(2009\)](#page-16-58). We retrieved the disk inclination from the VANDAM survey results [\(Segura-Cox et al. 2018\)](#page-16-46) and the position angle from the MASSES survey [\(Lee et al. 2016\)](#page-16-43). We only used the mass of the envelope determined in [Jørgensen et al.](#page-16-59) [\(2007\)](#page-16-59), $M_{\text{env}} = 2.3 \text{ M}_{\odot}$ for the total mass of the system, which is a good enough approximation to model the free-fall for a Class 0 protostar, as most of the mass is in the envelope. Table [3](#page-12-0) shows the set of parameters that can approximately reproduce the KDE of the observed HC_3N velocities within a region that surrounds the T_{MB} peak area. The resulting trajectory from these parameters is plotted in Fig. [7](#page-10-0) left. We note that this is a rudimentary adjustment of these velocities, and that the region we believe the streamer to cover has an equivalent size of two beams. At scales smaller than ∼ 1000 au from the protostar, the best free-fall solution diverges from the observed v_{LSR} distribution as there are possibly other gas components within the beam (for instance, protostellar envelope and disk rotation). Our current resolution limits our interpretation to scales larger than about 1500 au. Higher resolution observations of HC_3N line emission will allow for a better trajectory modelling.

The N_2H^+ velocity profile is significantly different than what HC_3N shows in this region. IRAS 4A is located at \approx 4" from a local N_2H^+ peak, as seen in Fig. [7](#page-10-0) top right. There is a velocity gradient from redshifted to blueshifted velocity with respect to the protostar's v_{LSR} (6.9 km s⁻¹) going from west to east (Fig. [7](#page-10-0) middle right). The eastern side of this gradient is blueshifted with respect to the protostar's v_{LSR} and is brighter than its redshifted counterpart. We made a position-velocity (PV) cut perpendicular to the outflow direction, shown in Fig. [7](#page-10-0) bottom right. The shape of the N_2H^+ PV diagram is suggestive of rotation within the envelope surrounding the protostar. IRAS 4A is expected to be within an envelope as it is classified as a Class 0 source. In previous works, gas surrounding IRAS 4A was found to be in-

Fig. 7: Zoom-in plots of HC₃N (left) and N₂H⁺ (right) emission for IRAS 4A, together with the best free-fall trajectory found for the T_{MB} and v_{LSR} maps. Top: Amplitude T_{MB} of the Gaussian components corresponding to HC₃N cluster no. 6 and to N₂H⁺ blue cluster. The black polygon marks the region selected as a potential streamer. The white curve marks the potential streamer's trajectory. The blue and red arrows indicate the direction of the outflow lobes for IRAS 4A and 4B, and the black arrow over the N_2H^+ map shows the orientation of the position-velocity diagram at the bottom. Middle: central velocities v_{LSR} of the Gaussian component. A scalebar representing 2000 au is at the top right of the image. Bottom left: KDE of the v_{LSR} within the selected region. The black line at the bottom of the plot represents a length of one beam. The blue curve marks the velocity versus the distance for the found free-fall solution. The black dashed line represents the protostar's v_{LSR} . Bottom right: N_2H^+ position-velocity diagram along the path indicated in the top panel. The dashed horizontal line marks the v_{LSR} of the source (6.9 km s⁻¹, [Stephens](#page-16-38) [et al. 2019\)](#page-16-38). The black scalebar represents a length equivalent to one beam.

falling [\(Belloche et al. 2006\)](#page-15-5). This profile then could be associated with a rotating-infalling envelope.

4.2.2. IRAS 4B

IRAS 4B is a Class 0 protostar that is to the southeast of IRAS 4A. Together with IRAS 4B2 they form a binary system. There is a bright HC_3N emission perpendicular to the outflow that dominates the emission around IRAS 4B, marked with a dashed black polygon in Fig. [8](#page-11-0) left. The direction of the lane coincides with the direction of the IRAS 4B2 outflow (Fig. [8](#page-11-0) left) and the velocity of this lane increases with increasing distance, so this bright lane is probably not asymmetric infall. However, at a 45 deg an-

gle from this lane, there is another signature of extended HC_3N emission, fitted with one Gaussian and belonging to cluster no. 6, and therefore part of the blueshifted fiber seen in Sect. [3.2.](#page-4-3) This smaller region is marked with a black polygon in Fig. [8](#page-11-0) top left. This bright emission has a velocity gradient that is consistent with infall toward IRAS 4B, based on the KDE of the observed velocities in Fig. [8](#page-11-0) bottom left (black KDE). The difference between the v_{LSR} of the protostar and the v_{LSR} of the candidate at the beginning of the streamer (at about 3000 au, Fig. [8\)](#page-11-0) is about 0.4 km s[−]¹ , which indicates that, to model this with a streamline model, we may require an initial radial velocity $v_{r,0} \neq 0$. Nevertheless, the central velocity of this protostar has been reported between 6.8 and 7.1 [\(Imai et al. 2018;](#page-16-40) [Stephens et al. 2019\)](#page-16-38). We

Fig. 8: Zoom-in plots of HC_3N emission for IRAS 4B (left), SK 15 (center) and IRAS 2B (right), with the same panels as shown in Fig. [7](#page-10-0) left. The Gaussian component for each protostar is labeled accordingly. The black polygon marks the region selected as a potential streamer. The black ellipse at the bottom left corner represents the beam. For IRAS 4B, the dashed black polygon represents the region analyzed in relation to IRAS 4B2. Top: Amplitude T_{MB} of the Gaussian component plotted. The blue and red arrows indicate the direction of the blueshifted and redshifted outflow lobes, respectively, for known outflows in the plotted area. Middle: Central velocity v_{LSR} of the Gaussian component selected. The scalebar represents a length of 2000 au. Bottom: KDE of the v_{LSR} within the selected region. The black density histograms represent the KDE of the velocities within the black polygons. The dashed lines mark the v_{LSR} of each protostar. The black scalebar represents a length equivalent to one beam. For IRAS 4B, the blue KDE represents the velocities within the dashed polygon.

chose the value from [\(Stephens et al. 2019\)](#page-16-38) as it is the one obtained with the best resolution.

At the position of IRAS 4B, there is a local N_2H^+ brightness peak. By doing a PV cut on the direction perpendicular to the outflow direction (Fig. [D.1](#page-20-0) left), we observed that this peak is slightly blueshfted with respect to the protostellar v_{LSR} . There is no sign of a rotating or infalling envelope in the N_2H^+ surrounding either IRAS 4B or 4B2: both regions are dominated by emission at fiber scales.

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4.2.3. SK 15

SK 15 (from the nomenclature of [Sandell & Knee 2001\)](#page-16-60) is a Class I protostar toward the southeast of SVS13A. Its v_{LSR} is estimated between 8 $km s^{-1}$ [\(Imai et al. 2018\)](#page-16-40) and 8.1 $km s^{-1}$ [\(Higuchi et al. 2018\)](#page-16-61). To the best of our knowledge, it does not have an observed outflow of its own, but lies close to the out-flow driven by SVS 13A (northern part of the map in Fig. [5\)](#page-7-1). We found a streamer candidate in HC_3N emission in the two-Gaussian fit around this protostar (Fig. [8](#page-11-0) center). One of the $HC₃N$ Gaussian components is part of cluster no. 4, within the redshifted fiber (Fig. [3\)](#page-5-2), whereas the other is recognized as noise by HDBSCAN. This extra component is the weakest of the two

Fig. 8 (cont.): Zoom-in plots of HC3N emission for SVS 13A (left), VLA 3 (center) and Per-emb 15 (right), with the same panels as shown in Fig. [7](#page-10-0) left. For VLA 3, the red dashed line marks the v_{LSR} of the fiber component toward the position of the protostar.

Table 3: Parameters of the trajectory that best reproduce the $HC₃N$ observations around IRAS 4A.

Parameter	Unit	Value
ϑ_0	deg	92
φ_0	deg	-5
r_0	au	3000
$V_{r,0}$	$km s^{-1}$	0
Ω_0	s^{-1}	2.3×10^{-13}
i^*	deg	-35
** PΔ	deg	19

Notes. (***) Angle obtained from the disk inclination [\(Segura-Cox et al.](#page-16-46) [2018\)](#page-16-46). (****) Angle obtained from the outflow direction [\(Lee et al. 2016\)](#page-16-43).

in T_{MB} (0.5 K versus 2.2 K). Looking at the component that is not part of the larger filament, the tail-shaped extension has a velocity gradient that gets more blueshifted with respect to SK 15's vLSR as gas gets closer to SK 15 (Fig. [8](#page-11-0) center).

SK 15 is located within 7" of a N_2H^+ T_{MB} peak. This peak is located at the position of the streamer candidate (Fig. [D.1](#page-20-0) middle). In this peak, there are two Gaussian components in the N_2H^+ spectral decomposition (Sect. [3.1\)](#page-3-2). The component responsible for the T_{MB} peak has a v_{LSR} of approximately 8.1 km s[−]¹ and belongs to the redshifted fiber. The other component, not clustered by HDBSCAN, has similar v_{LSR} to the HC₃N emission that we consider as streamer candidate, but it does not show a velocity gradient.

4.2.4. IRAS 2B

IRAS 2B (also known as Per-emb-36, [Enoch et al. 2009b\)](#page-16-62) is a Class I protostar located at the western edge of the ProPStar map. There is denser gas beyond the edge of the map, as shown in previous works covering this region in, for instance, $NH₃$ [\(Friesen](#page-16-63) [et al. 2017;](#page-16-63) [Dhabal et al. 2019\)](#page-16-31). Towards this edge, we found a bright extension of HC_3N gas that has a gradient from redshifted to blueshifted v_{LSR} with respect to IRAS 2B's v_{LSR} = 6.9 km s⁻¹

(Fig. [8](#page-11-0) right). This region is about the size of the HC_3N beam. We consider this as a streamer candidate as it shows a velocity gradient almost perpendicular to the outflow direction, although the complete gradient is contained within a beam. The brightness distribution suggests the streamer candidate could continue beyond the extend of the HC_3N map.

The N_2H^+ map also shows that the dense gas emission continues to the right of IRAS 2B. Previous N_2H^+ observations from [Dhabal et al.](#page-16-31) [\(2019\)](#page-16-31) show that there is considerable emission outside our coverage, toward the west of IRAS 2B, so we only get a glimpse of the kinematics toward the east of IRAS 2B. Figure [2](#page-4-2) shows that toward the northwest and southeast of the protostar, N_2H^+ emission can be fit with two Gaussians. We made a PV cut perpendicular to the outflow direction, shown in Fig. [D.1](#page-20-0) right. Within 4000 au from the protostar at each side, we see no clear indication of a rotating envelope, so we conclude that the N_2H^+ emission is dominated by the fiber kinematics.

4.2.5. SVS 13A

SVS 13A, also known as Per-emb-44 (using the nomenclature of [Enoch et al. 2009b\)](#page-16-62), is a Class 0/I close binary system (separated by approximately 70 au, [Anglada et al. 2000\)](#page-15-6). It drives powerful outflows and jets that dominate the CO emission of this region [\(Plunkett et al. 2013;](#page-16-29) [Stephens et al. 2019\)](#page-16-38). There is a detected streamer toward this source seen in DCN emis-sion by [Hsieh et al.](#page-16-57) [\(2023b\)](#page-16-57). The cut $HC₃N$ images toward SVS 13A show a beam-sized region consistent with the velocity gra-dient shown by [Hsieh et al.](#page-16-57) [\(2023b\)](#page-16-57). The HC_3N cluster no. 4 shows a sudden drop in v_{LSR} at the position of SVS 13A (Fig. [8](#page-11-0)) (cont.) left). The fiber's v_{LSR} around SVS 13A is between 8.4 and 8.5 km s^{-1} , and in the region where the streamer was found using DCN the velocity drops down to approximately 8.3 km s[−]¹ . This change is similar to the velocity gradient observed in the SVS 13A streamer [\(Hsieh et al. 2023b,](#page-16-57) in their Fig. 6). However, the HC_3N brightness distribution is difficult to interpret as this particular streamer. As seen in Fig. [8](#page-11-0) (cont.) top left, there is a local peak in T_{MB} located at 4.5″ east from the SVS 13A. This peak might blend with the less bright emission from the streamer. It is also possible that the streamer is small, about the size of our beam (~ 1500 au), and at this resolution, any HC₃N emission coming from the streamer is blended with the larger, fiber-dominated emission. Nevertheless, we were able to detect the small velocity gradient even in the blended image, which means that it is possible that $HC₃N$ emission traces the streamer toward SVS 13A. A higher spatial resolution image should be able to separate the emission corresponding to the fiber and the streamer.

SVS 13A lacks an N_2H^+ peak. The brightest peak near SVS 13A is located on SVS 13B to the southwest. We plot a PV diagram centered on SVS 13A and in the direction perpendicular to the outflow in Fig. [D.1](#page-20-0) (cont.) left. [Chen et al.](#page-16-64) [\(2009\)](#page-16-64) show that SVS 13B and VLA 3 are embedded in the same core, which shows signs of rotation. Therefore, to the west, the N_2H^+ velocity field is dominated by the SVS 13B and VLA 3 joint core. To the east, the velocity field is dominated by the extended N_2H^+ component tracing the redshifted fiber.

4.2.6. VLA 3

VLA 3 is a protostellar source located 10′′toward the west of SVS 13A. It does not have an outflow direction on record, nor information about a potential disk. However, it lies on the path

of the SVS 13B outflow. As mentioned in Sect. [4.2.5,](#page-13-0) VLA 3 and SVS 13B are embedded in a common core, suggesting these two sources are forming a bound binary system [\(Chen et al. 2009\)](#page-16-64). As with the case of SK 15, one of the HC_3N Gaussian components located near VLA 3 is part of the redshifted fiber, within the HDBSCAN cluster no. 4. The other $HC₃N$ Gaussian component is categorized as noise by the clustering algorithm. Looking at this remaining component, there is a velocity gradient that points toward VLA 3. We show the T_{MB} map and the observed velocities in the region of the streamer candidate in Fig. [8](#page-11-0) (cont.) center. As there is no recorded velocity for this protostar, we used the N_2H^+ fiber velocity at the location of VLA $3 (v_{LSR} = 8.4 \text{ km s}^{-1})$ as reference. The difference between the $v_{LSR} = 8.4 \text{ km s}^{-1}$ as and the N.H⁺ v_{LSR} decreases with increas v_{LSR} of the HC₃N gas and the N₂H⁺ v_{LSR} decreases with increasing distance from VLA 3. We also plot the outflow directions of SVS 13A and 13B to see if they interfere with the HC_3N gas that could potentially trace a streamer. The selected HC_3N component is not affected by the outflow: the redshifted cones of SVS 13A and 13B have $v_{LSR} > 10 \text{ km s}^{-1}$, whereas the HC₃N emis-
sion from this Gaussian component is lower than 8 km s⁻¹; the sion from this Gaussian component is lower than 8 km s^{-1} ; the $HC₃N$ velocity dispersion is average for its surrounding cloud (0.2 km s^{-1}) ; and finally, there is no discernible velocity gradient in the direction of the outflows (northwest).

Also similar to SK 15, there is a local N_2H^+ peak at the position of the streamer candidate to the north of VLA 3, with two Gaussian components (Fig. [D.1](#page-20-0) middle). Unlike the case of SK 15, the region where there are two Gaussian components is smaller than the area covered by the $HC₃N$ streamer candidate, barely the size of the N_2H^+ beam. The Gaussian component responsible for the peak is within the cluster that makes up the redshifted fiber and has a velocity of about 8.51 km s⁻¹. The second, weakest component shows an almost constant velocity between 7.88 and 7.91 km s[−]¹ , similar to the velocity of the streamer but without a clear gradient.

4.2.7. Per-emb 15

Per-emb-15 (also known as SK 14, [Sandell & Knee 2001;](#page-16-60) [Enoch](#page-16-62) [et al. 2009b\)](#page-16-62) is a Class I protostar located toward the south of the SVS 13 system. Within the HC_3N single Gaussian fit, there is a velocity gradient within a local T_{MB} peak (Fig. [8](#page-11-0) (cont.) right). This structure was not recognized in the HDBSCAN as part of any cluster, possibly because it is located where the fibers overlap. An inspection of the HC_3N spectra confirmed that there is a well-fitted single Gaussian component along the defined region, and not two Gaussian components fitted as one. The HC_3N velocity gradient follows the expected direction for asymmetric infall toward the protostellar disk (Fig. [8](#page-11-0) (cont.) bottom right).

Per-emb 15 has no N_2H^+ peak within a beam of its location. the closest N_2H^+ brightness peak is at about 18″ to the northwest, where the outflow traced by MASSES seems to end. The N_2H^+ emission of cluster no. 2 shows a blueshift in velocity toward the protostar. This is more clear in Fig. [D.1](#page-20-0) (cont.) right. We made a PV diagram perpendicular to the outflow direction, shown in Fig. [D.1](#page-20-0) (cont.) right bottom. This diagram is dominated by the emission from the fibers, as this source is located where the two fibers overlap in our line of sight. The blueshifted emission might be produced by the protostar or might be part of the velocity gradient seen in the fiber and caused by gas infall toward the fiber spine.

5. Discussion

5.1. Infall of gas onto the fibers

Our results suggest that HC_3N and N_2H^+ trace different structures toward NGC 1333 SE. First, both molecules follow the filament direction, but their distributions are different. N_2H^+ recovers more area within the fibers located in this region (Sect. [3.1\)](#page-3-2), whereas HC_3N is more patchy and is affected by the presence of outflows (Sect. [4.1.2\)](#page-5-1). Part of the difference in their structure can be due to the different densities traced, where $HC_3N (10-9)$ molecular line has a critical density about an order of magnitude higher than the N₂H⁺ (1 – 0) line, as well as a higher E_{up} (Table [1\)](#page-3-1). HC_3N has a higher T_{peak} closer to the protostars. However, we also detected a systematic redshift of HC_3N with respect to N_2H^+ in the case of the redshifted fiber, as shown in Sect. [4.1.3.](#page-7-0) This result shows that the kinematics of the HC_3N gas are different from N_2H^+ gas in the redshifted fiber, and thus, the HC_3N is not part of the dense fiber structure but traces dense gas that is less chemically evolved than the gas in the fibers.

For the blueshifted fiber, the difference along the line of sight is not significant when compared to the velocity uncertainties (Sect. [4.1.3\)](#page-7-0). However, it is possible that this is a projection effect and that the motion of \overline{HC}_3N with respect to N_2H^{\dagger} is mostly perpendicular to our line of sight. Previous works have suggested that gas is being pushed by an expanding bubble, responsible for the formation of the filament, from the southwest of IRAS 4A [\(Dhabal et al. 2019;](#page-16-31) [De Simone et al. 2022\)](#page-16-65). $HC₃N$ emission traces the blueshifted fiber only toward the west side, which fits into this picture if this newly deposited gas is moved by this expanding shell. We note that close to IRAS 4A and $4B$, $HC₃N$ gas is considerably blueshifted with respect to N_2H^+ gas. This is due to the effect of the outflows of these two protostars that stir the fresh gas more than the dense core structure (traced by N_2H^+).

Given the different kinematics and structures traced, we propose that the HC_3N emission represents a layer of chemically fresh material feeding the fibers. In chemical models of lowmass star forming regions, HC_3N appears at earlier times (a few 100 kyr from $t = 0$) than nitrogen-bearing molecules like N_2H^+ [\(Bergin & Tafalla 2007;](#page-15-4) [Sakai & Yamamoto 2013\)](#page-16-66), as nitrogen chemistry starts from neutral-neutral reactions, which are much slower than the ion-neutral reactions that dominate the carbon chemistry. HC_3N has also been suggested as a product of outflow shock fronts [\(Shimajiri et al. 2015\)](#page-16-67), and in our work we observe that outflows help to stir the gas traced in HC_3N , but we do not include the velocity components of HC_3N with outflow wings in our clustering (Sect. [3.2\)](#page-4-3) and subsequent structures' analysis. This ensures that the HC₃N gas that we compare with N_2H^+ is not affected by possible shock chemical enhancement. This indicates that there is chemical replenishment, delivered from the patchy structure seen in HC_3N . The fact that HC_3N is detected in a smaller extension and is patchy with respect to N_2H^+ suggests that the region could be more chemically evolved than other regions in Perseus, e.g. in Barnard 5, where HC_3N is more extended than the filaments' structures traced by $NH₃$ [\(Valdivia-](#page-16-13)[Mena et al. 2023\)](#page-16-13). Nevertheless, there is still some fresh material for star formation, in the shape of sparsely distributed gas.

The presence of N_2H^+ perpendicular gradients in both fibers suggests that gas falls toward the fiber spines. These are most prominent toward the south of the filament, where there are no protostars. These kind of gradients are also described for this region in [Dhabal et al.](#page-16-32) [\(2018\)](#page-16-32) and [Chen et al.](#page-16-33) [\(2020b\)](#page-16-33). [Dhabal](#page-16-32) [et al.](#page-16-32) [\(2018\)](#page-16-32) suggest the presence of a global velocity gradient for what they call subfilament A, which corresponds in velocity to our redshifted fiber, but it is unclear for their subfilament B,

corresponding approximately in v_{LSR} with our blueshifted fiber. We confirmed that there is also a perpendicular velocity gradient for the blueshifted fiber in N_2H^+ . As suggested in these previous works, the velocity gradients observed in the fibers can be related to accretion flow toward the fibers, seen as contraction of a sheetlike cloud [\(Chen et al. 2020a\)](#page-16-68). The velocity gradients could also be consistent with fiber rotation, but in simulations of filament formation, these motions do not seem favorable in fibers [\(Smith](#page-16-69) [et al. 2016\)](#page-16-69). HC_3N also shows prominent perpendicular velocity gradients where there are no protostars, but the regions which follow this motion are smaller. In general, the gradients in HC_3N tend to change more drastically near protostars (Fig. [4](#page-6-0) top). This suggests that the presence of the protostars stirs the surrounding gas, which is reflected in more complex gas motions at smaller scales.

5.2. Discovery of streamer candidates in NGC 1333 SE

We found streamer candidates toward 7 out of 16 YSOs within our field of view. Moreover, we confirmed the infalling motion toward one of them (IRAS 4A). For the rest of the candidates, we do not have enough spatial resolution or information about the protostar and disk masses to efficiently model the free-falling gas. What we can say about their three-dimensional structure is that these streamers come from behind the protostars, as they all become more blueshifted with respect to the protostellar v_{LSR} with decreasing distance. The streamer candidates in our work were found towards Class 0, 0/I and I protostars. This represents a streamer frequency of approximately 40% if we consider all the YSOs within the maps' field of view. If we consider only the early stage protostars (Class 0 to I, for which we have 12 within the map), the frequency of streamers is higher, about 60%. Although these are small number statistics, it is a first approach to quantify the prevalence of streamers toward YSOs (if we are able to confirm the infall nature of the emission in the future).

This is a first coarse estimate to the real frequency of streamers in the area, as these are candidates still not confirmed dynamically, and it is possible we have missed streamers. First, we are unable to observe any streamer that is smaller than approximately 1500 au in projected length, as that is our resolution limit. This decreases our chances to find small streamers and also those almost fully contained along our line of sight. In particular, this limits our chances to find Class II streamers (of which we do not find any), as they are usually on the order of 300 to 500 au [\(Garufi et al. 2022;](#page-16-9) [Ginski et al. 2021\)](#page-16-15), although longer arms of up to ∼ 2000 au surrounding T Tauri disks have been suggested to be streamers [\(Alves et al. 2020;](#page-15-7) [Huang et al. 2021,](#page-16-70) [2022\)](#page-16-71). Most Class II sources in the field have no detected HC_3N emission: only one Class II source (ASR 54) has $HC₃N$ emission detected around it, and its velocity structure is not suggestive of streamer motion, but appears dominated by the infall of fresh gas into the fiber. This does not mean that there are no streamers toward these sources, but the tracer (HC_3N) and/or the resolution are not adequate to find streamers toward these sources. Secondly, the assumption of monotonically decreasing velocity with distance along the line of sight can miss streamers. This assumption helps to recognize infalling motion and to differentiate it from outflowing gas, but there are instances where a streamer will not show this kind of gradient along the line of sight. If the acceleration is completely contained in the plane of the sky, the observed v_{LSR} of the gas is constant. It is also possible that, due to projection effects, the velocity appears to be monotonically *increasing* in the outer regions of the streamer, such as in the case of Per-emb 2 [\(Pineda et al. 2020\)](#page-16-10), confirmed using the free-fall

analytic solution from [Mendoza et al.](#page-16-58) [\(2009\)](#page-16-58). As an acceleration proportional to distance is usually attributed to outflowing gas, it is necessary to take other factors into account (like the shape of the emission and the position of the known outflow) to disentangle outflowing emission from inflowing gas. Considering the spatial and spectral resolution of our HC_3N data and the median velocity gradient found in our streamers (~ 15 km s⁻¹ pc⁻¹), the probability to miss a streamer is about 50%, either because it is contained along our line of sight or close enough to the plane of the sky that its gradient is not observable.

We define the structures found in this work, including toward IRAS 4A, as streamer candidates. The limiting factors at the time of describing the structure of the streamers is the lack of angular resolution and the lack of information about the protostar and disk masses. We show that HC_3N traces infall given the velocity of the gas that surrounds a protostar and the location of the local velocity gradient, but we do not have enough information about the protostars themselves to confirm that the velocity profiles correspond to free-fall. To confirm the streamer nature of these candidates, it is necessary to replicate their structure (a thin and long structure in T_{MB} and a velocity gradient) with a free-fall model. In the case of IRAS 4A, even if we could model the infall, the resolution is not high enough to describe the structure in the image plane. Therefore, we require follow-up observations with a resolution higher than 4.9′′to fully describe the structure of these streamers.

5.3. Relation of streamers to the larger gas infall

Our results indicate that the gas that builds up the streamers comes from beyond the fibers. There are three main reasons for this conclusion. First, we did not observe signs of infall using N_2H^+ . The N_2H^+ v_{LSR} close to each protostar is different than the HC_3N v_{LSR}, and we do not observe stream-like structures in N_2H^+ , although this could be due to projection effects. Second, by construction, the structure of the fibers is traced using N_2H^+ , and we observed significant differences in the central velocities of the redshifted fiber with respect to the $HC₃N$ flows within. In the case of the blueshifted fiber, we did not see a significant difference (Sect. [4.1.3\)](#page-7-0). However, this might be a projection effect, where we did not see signs of any kinematic difference because the movement is along the plane of the sky. Third, the apparent direction of the streamer candidates does not coincide with the orientation of the filament.

Our results suggest that the mass that composes the streamers does not come from the fiber structure, but from the fresh gas that is infalling toward the fibers. Although $HC₃N$ can be enhanced by the presence of outflows, we took out the velocity components affected by outflows for this analysis, so the streamer gas being chemically fresh is the most probable origin. A similar relation between gas outside the filament and streamers is suggested for Barnard 5, another region in Perseus, but with different tracer for each scale [\(Valdivia-Mena et al. 2023\)](#page-16-13). This result is consistent with research that indicates that cores must be replenished with fresh gas to form protostars, as the amount of envelope mass dispersed through the outflow is substantial [\(Hsieh et al. 2023a\)](#page-16-72). Streamers, therefore, can be the mechanism that can feed the protostellar system with mass that is not from the original core, connecting the material coming from outside the fibers and delivering it toward the protoplanetary disks.

6. Conclusions

In this work, we analyzed the distribution and velocity structure of the NGC 1333 SE fibers in search for accretion signatures toward YSOs. Our results are summarized below.

- The distribution of HC_3N gas in this region is patchy and does not cover the full extent of the fibers seen in N_2H^+ . The HC_3N velocity along the line of sight is redshifted with respect to N_2H^+ in the redshifted fiber. Together, these results indicate that HC_3N is following different kinematics than N_2H^+ . We suggest HC_3N traces gas that is infalling later to the filament, after star formation has started in the region.
- N_2H^+ shows velocity gradients perpendicular to the fibers' orientations. We suggests, as previous works have done as well [\(Chen et al. 2020b;](#page-16-33) [Dhabal et al. 2018\)](#page-16-32), that this indicates infall toward the fibers's spines.
- The outflows of IRAS 4A and IRAS 2A generate wings and strong brightness peaks in the HC_3N spectra. In these regions, HC_3N is enhanced at the bow shocks of the protostellar outflows.
- We found streamer candidates toward 7 out of 16 YSOs in the field of view of our mosaic. This represents an incidence of about 40% of YSOs with streamers when looking within a region. The 7 candidates are all found toward early stage protostars (Class 0 to I), which represent a total of 12 sources within our field of view, so for early stages in particular, the incidence of streamers is about 60% for this given filament.
- The gas that composes the streamers is coming from outside the fiber structure, as there is a difference between the velocity structure of HC_3N with N_2H^+ and no streamers are detected in N₂H⁺ emission. Only 2 protostars show N₂H⁺ emission with similar velocities in the same position as the $HC₃N$ streamers, but do not show the same velocity profiles.

We defined these structures as candidate streamers because the resolution of our data is not enough to resolve the width of the flow and does not allow for an accurate modelling of the infall. Further information on the mass and orientation of the protostars and disks in the region are required to model the infall and determine the true length of the streamers. If we take NGC 1333 as a "typical" star-forming region, then we expect streamers to be a frequent feature toward protostellar envelopes. This work highlights the relevance of streamers in our new picture of low-mass star formation.

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Appendix A: Channel maps

Figures [A.1](#page-18-0) and [A.2](#page-18-1) show the channel maps of HC_3N and N_2H^+ emission between approximately 6 and 9.2 km s^{-1} , in steps of about 0.35 km s⁻¹. We note that there is emission in HC₃N outside of this range, starting from -3 km s⁻¹ and ending at approximately 13 km s^{-1} .

Appendix B: Density-based clustering of molecular emission

In this section, we describe the steps taken to cluster the N_2H^+ and HC_3N velocity structures. We cluster the Gaussian points from each molecule based on their position in the plane of the sky and their v_{LSR} . Using the amplitude T_{peak} or their dispersion σ_{v} does not improve the clustering results, and as our goal is to disentangle the velocity structures of the filament, the analysis in the position-position-velocity space is sufficient (in particular after filtering out the Gaussians with high dispersion).

From the Gaussian fitting and the quality assessment, we have 94864 Gaussians in the N_2H^+ data cube and 56771 Gaussians in the HC_3N data cube. We did a finer selection of Gaussian components for both molecules before clustering, so as to ensure that we trace the bulk emission of the filament structure. We select only the points with low uncertainty in their parameters, that is, with an uncertainty less than 25% the value of the parameter. This criterion leaves out uncertain points that can add confusion to the clustering. We also filter out the Gaussian fits that show $\sigma_v > 1$ km s⁻¹, as these represent fits with outflows. We filter possible Gaussian components that represent random noise in possible Gaussian components that represent random noise in the spectra by selecting only those components with velocities between 5 and 9 km s⁻¹. This also helps to filter out HC₃N that corresponds to high-velocity outflow wings. We were left with 81657 Gaussians in the N_2H^+ data cube and 52019 Gaussians in the HC_3N data cube.

We cluster the Gaussian component results of N_2H^+ using Hierarchical Density-Based Spatial Clustering of Applications with Noise (HDBSCAN). HDBSCAN is an extension of the Density-Based Spatial Clustering of Applications with Noise (DBSCAN) algorithm. In summary, DBSCAN defines clusters of points in the user-defined hyperspace as local point overdensities, and leaves sparsely distributed points as noise [\(Ester et al.](#page-16-73) [1996\)](#page-16-73). A core is defined as an overdensity surrounding a core sample (a single point in the hyperspace) with a minimum of samples *n* within a radius of ϵ . HDBSCAN, instead of fixing the radius ϵ , selects clusters based on the minimum spanning tree of the mutual reachability graph^{[5](#page-17-3)}, i.e. explores all possible values of ϵ [\(Campello et al. 2013;](#page-16-74) [McInnes & Healy 2017\)](#page-16-75). This process allows the algorithm to form clusters of different densities. We use the hdbscan package from the contributed packages to scikit-learn 6 6 6 .

Table B.1: HDBSCAN parameters used to cluster the HC_3N and N_2H^+ Gaussian peaks.

Parameter	HC, N	$N_{\alpha}H^{+}$
min_sample	200	240
min cluster size	500	900

The best results for each molecule are obtained with the parameters shown in Table [B.1,](#page-17-5) named according to the parameter

names in the python implementation^{[7](#page-17-6)}. The parameters are different for each molecule because the number of Gaussian peaks for $HC₃N$ (about 51000) is much smaller than the number of Gaussian peaks in N_2H^+ (more than 81000). The resulting clusters can be seen in Fig. [3.](#page-5-2)

Figure [B.1](#page-19-0) shows the peak temperature T_{peak} , velocity v_{LSR} and dispersion $\sigma_{\rm v}$ of the N_2H^+ clusters, grouped according to their fiber redshifted or blueshifted resulting from the HDBtheir fiber, redshifted or blueshifted. resulting from the HDB-SCAN analysis. Figure [B.2](#page-19-1) shows the same quantities but for the N_2H^+ clusters.

Appendix C: Comparison of clustering results with previous works

Figure [3](#page-5-2) show that the positions of the clusters found for N_2H^+ emission are consistent with the $NH₃$ fibers found in [Chen et al.](#page-16-33) [\(2020b\)](#page-16-33). However, our N_2H^+ clustering is different than the fiber structure found by [Hacar et al.](#page-16-2) [\(2017\)](#page-16-2), using the same molecule and hyperfine transition. Figure [C.1](#page-20-1) left shows the N_2H^+ fiber spines from [Hacar et al.](#page-16-2) [\(2017\)](#page-16-2) over the clusters found in N_2H^+ emission. Our clustering does not recognize a third separate fiber toward the SVS 13 region as in the case of [Hacar et al.](#page-16-2) [\(2017\)](#page-16-2). The redshifted cluster found in our work coincides well with the leftmost fiber from [Hacar et al.](#page-16-2) [\(2017\)](#page-16-2) (as well as the eastern fiber from [Chen et al. 2020b,](#page-16-33) located at the same position). The blueshifted cluster, on the other hand, does not coincide with the rightmost fiber from [Hacar et al.](#page-16-2) [\(2017\)](#page-16-2), which connects better to the southernmost part of our redshifted cluster. This mismatch between structures found using the same molecule and transition is due to the different algorithms used for structure recognition and the differences in resolution between the datasets.

We do not expect to obtain the same fiber structures as previous works as we use a different, more general algorithm. HDB-SCAN is a purely mathematical algorithm with no physical information. Our clustering is based on proximity in a scaled PPV space, without considering other parameters such as brightness or line width, whereas Friends in Velocity (FIVE) is designed to identify velocity-coherent structures based on emission intensity and velocity, considering the PPV space is Nyquist-sampled and represents a fluid [\(Hacar et al. 2013\)](#page-16-76).

The different resolutions contribute to the different structures recognized by the different algorithms. The beam of our data (∼ 6 ′′, Table [1\)](#page-3-1) is about 5 times smaller than for the data used in [Hacar et al.](#page-16-2) [\(2017\)](#page-16-2) (30 ′′), so our data resolves the structure *within* the fibers identified in previous works. When the beam increases, the brightness sensitivity is better, but the power to resolve substructures decreases, so even though the emission that corresponds to our blueshifted fiber might have been detected with IRAM 30-m, it was not large enough to categorize in its own structure and was left as noise by FIVE. Our clustering is used as a guide to match which velocity components from different molecules can be studied together, so as to understand the difference in kinematics between N_2H^+ and HC_3N (Sect. [4.1.3\)](#page-7-0). Also, at this scale, fiber-identification algorithms such as FIVE will probably find further substructures. Future works will explore what kind of substructures can be found using fiber-identification algorithms in high-resolution observations. The comparison of our clustering results with their recovered fibers shows that structure identification in PPV data is dependent on both the data itself and the algorithm chosen to identify structure.

⁵ https://scikit-learn.org/stable/modules/clustering.html

⁶ https://hdbscan.readthedocs.io/en/latest/

⁷ https://hdbscan.readthedocs.io/en/latest/parameter_selection.html

Fig. A.1: Channel maps between 6.02 and 9.17 km s⁻¹ for HC₃N $J = 10 - 9$ emission.

Fig. A.2: Channel maps between 6.02 and 9.17 km s⁻¹ for N_2H^+ *J* = 1 − 0, $F_1F = 01 - 12$ emission.

Appendix D: Close-up of N² H ⁺ velocity profiles toward individual protostars

Figure [D.1](#page-20-0) shows the same two top panels (T_{peak} and v_{LSR}) of Fig. [7](#page-10-0) right for all other sources where we find a streamer candidate. Some show position-velocity cuts, and for those where a second component in N_2H^+ coincides roughly with the streamer's velocity, we plot the v_{LSR} KDE of the corresponding component instead of the position-velocity cut.

Fig. B.1: Peak temperature, central velocity and velocity dispersion of the N₂H⁺ clusters. Top: T_{peak} , v_{LSR} and σ_{v} for the redshifted fiber. The redshifted fiber and σ_{v} for the blueshifted fiber. The red fiber. The blue contour indicates the position of the blueshifted fiber. Bottom: T_{peak} , v_{LSR} and σ_{v} for the blueshifted fiber. The red contour indicates the position of the redshifted fiber.

Fig. B.2: Same as Fig. [B.1](#page-19-0) but for the HC_3N clusters.

Fig. C.1: Clusters corresponding to the redshifted and blueshifted fibers in N_2H^+ and HC_3N (as Fig. [3](#page-5-2) left), with the N_2H^+ fiber spines taken from [Hacar et al.](#page-16-2) [\(2017\)](#page-16-2), shown as dashed black lines.

Fig. D.1: Zoom-in plots of N_2H^+ emission for IRAS 4B (left), SK 15 (center) and IRAS 2B (right), with the same panels as shown in Fig. [7](#page-10-0) right except for SK 15. The Gaussian component for each protostar is labeled accordingly. The black ellipse at the bottom left corner represents the beam. Top: Amplitude T_{MB} of the Gaussian component plotted. The blue and red arrows indicate the direction of the blueshifted and redshifted outflow lobes, respectively, for known outflows in the plotted area. Middle: Central velocity v_{LSR} of the Gaussian component selected. The scalebar represents a length of 2000 au. Bottom: For IRAS 4B and IRAs 2B, the positionvelocity diagram along the path indicated in the top panel. For SK 15, KDE of the v_{LSR} within the selected region. The red density histogram representd the KDE of the velocities within the black polygon. The dashed lines mark the v_{LSR} of each protostar. The black scalebar represents a length equivalent to one beam.

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Fig. D.1 (cont.): Zoom-in plots of N_2H^+ emission for SVS 13A (left), VLA 3 (center) and Per-emb 15 (right), with the same panels as shown in Fig. [7,](#page-10-0) except for VLA 3.