

# Bounds on new neutrino interactions from the first $\text{CE}\nu\text{NS}$ data at direct detection experiments

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Recently, two dark matter direct detection experiments have announced the first indications of nuclear recoils from solar  $^8\text{B}$  neutrinos via coherent elastic neutrino-nucleus scattering ( $\text{CE}\nu\text{NS}$ ) with xenon nuclei. These results constitute a turning point, not only for dark matter searches that are now entering the *neutrino fog*, but they also bring out new opportunities to exploit dark matter facilities as neutrino detectors. We investigate the implications of recent data from the PandaX-4T and XENONnT experiments on both Standard Model physics and new neutrino interactions. We first extract information on the weak mixing angle at low momentum transfer. Then, following a phenomenological approach, we consider Lorentz-invariant interactions (scalar, vector, axial-vector, and tensor) between neutrinos, quarks and charged leptons. Furthermore, we study the  $U(1)_{\text{B-L}}$  scenario as a concrete example of a new anomaly-free vector interaction. We find that despite the low statistics of these first experimental results, the inferred bounds are in some cases already competitive. For the scope of this work we also compute new bounds on some of the interactions using  $\text{CE}\nu\text{NS}$  data from COHERENT and electron recoil data from XENONnT, LUX-ZEPLIN, PandaX-4T, and TEXONO. It seems clear that while direct detection experiments continue to take data, more precise measurements will be available, thus allowing to test new neutrino interactions at the same level or even improving over dedicated neutrino facilities.

Keywords: new interactions, dark matter detectors, solar neutrinos,  $\text{CE}\nu\text{NS}$

## I. INTRODUCTION

Coherent elastic neutrino-nucleus scattering ( $\text{CE}\nu\text{NS}$ ) is a neutral-current process in which a low-energy neutrino scatters off an entire nucleus [1]. Its first theoretical prediction and the principles for its detection date back to the 1970s and 1980s [2, 3]. One main feature of  $\text{CE}\nu\text{NS}$  is that its Standard Model (SM) cross section is large compared to other neutrino scattering cross sections, as it is coherently enhanced being proportional to the number of nucleons squared. On the other hand, its experimental detection faces technological difficulties, as it requires the observation of nuclear recoils of very low energy. For this reason, this process evaded detection during many decades until its first observation by the COHERENT Collaboration [4], using an intense spallation source producing neutrinos from pions decaying at rest. Further observations and evidence using different targets [5–7] or a reactor source [8] have provided valuable information on the  $\text{CE}\nu\text{NS}$  cross section and its dependence on  $N^2$ .

The possibility that  $\text{CE}\nu\text{NS}$  detectors could be used as dark matter (DM) [9] detectors was pointed out by Goodman and Witten [10], anticipating the same experimental challenges. An intense experimental program [11, 12] has followed this early suggestion, culminating in current ton-scale DM direct detection experiments. The latest generation of low-threshold dual-phase liquid xenon (LXe) detectors, including the XENONnT [13], LUX-ZEPLIN (LZ) [14], and PandaX-4T [15]

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experiments, has reached impressive sensitivities, although without indicating any conclusive evidence of DM detection. As a by-product, the increase in target size has allowed these facilities to be sensitive to fluxes of astrophysical neutrinos. As anticipated, the improvement in the energy threshold at these experiments has now led to detectable rates of  $^8\text{B}$  solar neutrinos inducing CE $\nu$ NS [16–18]. Neutrino backgrounds from natural [19] and artificial [20] sources do pose a challenge for DM searches in the form of a *neutrino fog* [21], but at the same time they provide new opportunities to probe the neutrino sector [22–40].

A steady increase in sensitivity has allowed the XENONnT [41] and PandaX-4T [42] Collaborations to recently report their first indications of nuclear recoils from solar  $^8\text{B}$  neutrinos via CE $\nu$ NS. With the present exposures reached in these experiments, the background-only hypothesis is disfavored with a statistical significance of  $2.73\sigma$  in the case of XENONnT and  $2.64\sigma$  for PandaX-4T. These results update previous searches by the same Collaborations [43, 44] and constitute the first indication of nuclear recoils from solar neutrinos and the first CE $\nu$ NS observation on a xenon target. Assuming that no new physics is present, these results also provide a measurement of the  $^8\text{B}$  solar neutrino flux component which is in agreement with theoretical predictions [45] and with results from dedicated neutrino experiments [46–49]. These successful experimental results have immediately motivated new phenomenological studies, aiming at probing new physics in the form of non-standard neutrino interactions (NSI) [50, 51], light mediators [52] and also the determination of the weak mixing angle at low momentum transfer [53].

In this paper, we study some implications of the first CE $\nu$ NS indications at DM experiments both for SM and new physics. Following Ref. [53], we first revisit the determination of the weak mixing angle at the  $\mathcal{O}(10)$  MeV scale, providing also a result in terms of a combined (XENONnT and PandaX-4T) analysis. Next, we confront new interactions between neutrinos, quarks and charged leptons with the recent XENONnT and PandaX-4T data. We focus on Lorentz-invariant interactions (scalar, vector, axial-vector and tensor), parameterized in a model-independent way in the form of neutrino generalized interactions (NGIs) [54–58]. We consider both regimes of effective operators and light mediators, since direct detection experiments have low-energy thresholds and are hence sensitive to interactions involving light mediators [26, 59–61]. In addition, for the sake of example, we consider a motivated, anomaly-free  $U(1)$  extension of the SM with a B-L symmetry (B being the baryon number and L the total lepton number).

The remainder of this paper is organized as follows. In Sec. II we introduce the relevant CE $\nu$ NS cross sections, both in the SM and in the presence of NGIs. We discuss in Sec. III the simulation details as well as the procedure followed for the statistical analysis of XENONnT and PandaX-4T data. We present in Sec. IV our results in terms of a determination of the weak mixing angle at low energy and exclusion limits on the NGI parameter space. Finally, we draw our conclusions in Sec. V.

## II. COHERENT ELASTIC NEUTRINO-NUCLEUS SCATTERING CROSS SECTIONS

In this section, we provide the relevant CE $\nu$ NS cross sections, in the SM and in the presence of NGIs, required for the calculation of the corresponding event rates.

### A. CEνNS cross section in the Standard Model

In the SM, the CEνNS differential cross section with respect to the nuclear recoil energy  $T_N$ , neglecting  $T_N/m_N$  and higher order  $\mathcal{O}(T_N^2)$  terms, reads [2, 62]

$$\left. \frac{d\sigma_{\nu N}}{dT_N} \right|_{\text{SM}} = \frac{G_F^2 m_N}{\pi} (Q_V^{\text{SM}})^2 F_W^2(|\mathbf{q}|^2) \left( 1 - \frac{m_N T_N}{2E_\nu^2} - \frac{T_N}{E_\nu} \right), \quad (1)$$

with  $G_F$  being the Fermi constant,  $E_\nu$  the incoming neutrino energy, while  $m_N$  is the nuclear mass and  $Q_V^{\text{SM}}$  denotes the SM weak charge which is given by

$$Q_V^{\text{SM}} = g_V^p Z + g_V^n N, \quad (2)$$

where  $Z$  ( $N$ ) is the proton (neutron) number, and the proton and neutron couplings (at tree level)<sup>1</sup> are given by  $g_V^p = (1 - 4\sin^2\theta_W)/2$  and  $g_V^n = -1/2$ , respectively. The weak charge encodes the dependence on the weak mixing angle  $\theta_W$  through the proton contribution. From RGE extrapolation, its value in the low-energy regime is expected to be  $\sin^2\theta_W = 0.23857(5)$  [64]. Nuclear physics corrections are included in the form factor  $F_W^2(|\mathbf{q}|^2)$ , to account for the finite nuclear spatial distribution. Given the small momentum transfer involved in the CEνNS of <sup>8</sup>B solar neutrinos, the dependence on the form factor is small. At the scope of the numerical calculations, we rely on the Klein-Nystrand parametrization [65]

$$F_W(|\mathbf{q}|^2) = 3 \frac{j_1(|\mathbf{q}| R_A)}{|\mathbf{q}| R_A} \left( \frac{1}{1 + |\mathbf{q}|^2 a_k^2} \right), \quad (3)$$

where  $j_1(x) = \sin(x)/x^2 - \cos(x)/x$  is the spherical Bessel function of order one,  $a_k = 0.7$  fm and  $R_A = 1.23 A^{1/3}$  indicates the root mean square (RMS) radius (in [fm]),  $A$  being the atomic mass number. The expected magnitude of the momentum transfer is  $|\mathbf{q}| = \frac{\sqrt{2m_N T_N}}{197.327} \text{ fm}^{-1} \sim \mathcal{O}(10)$  MeV.

### B. CEνNS cross section with neutrino generalized interactions

One of our goals in the present work is to explore the implications of the recent XENONnT and PandaX-4T data on new neutrino interactions. For simplicity, we adopt a phenomenological approach and consider all possible Lorentz-invariant low-energy neutral-current interactions parameterized through the following effective Lagrangian

$$\mathcal{L}_{\text{NC}}^{\text{NGI}} \supset \frac{G_F}{\sqrt{2}} \sum_{\substack{a=(S,P,V,A,T), \\ \ell=e,\mu,\tau}} C_a (\bar{\nu}_\ell \Gamma^a P_L \nu_\ell) (\bar{N} \Gamma_a N), \quad (4)$$

where  $\Gamma^a = \{\mathbb{I}, i\gamma^5, \gamma^\mu, \gamma^\mu \gamma^5, \sigma^{\mu\nu}\}$  (with  $\sigma^{\mu\nu} = i[\gamma^\mu, \gamma^\nu]/2$ ),  $P_L \equiv (1 - \gamma^5)/2$  is the left-handed projector and  $N$  denotes the nucleus. The  $C_a$  are dimensionless coefficients which denote the corresponding neutrino-nucleus couplings for all interactions: scalar ( $S$ ), pseudoscalar ( $P$ ), vector ( $V$ ), axial-vector ( $A$ ) and tensor ( $T$ ). Notice that we consider only flavor-independent interactions and hence assume that the coupling  $C_a$  is the same for each neutrino flavor. Therefore, in the present analysis we do not need to include neutrino oscillations, which—in addition to the interactions—would also be modified by flavor-dependent interactions due to matter effects in the Sun, relevant for the energies typical of <sup>8</sup>B neutrinos.

<sup>1</sup> At higher orders these factors become flavor-dependent; the correction to  $g_V^n$  is very small, while the correction to  $g_V^p$  is quite significant [63], although the proton coupling remains very small in comparison to its neutron counter part.

Because of the typical momentum transfer involved in the CEvNS of solar neutrinos, we aim to extend our phenomenological study to the case of interactions involving light mediators (i.e., with a mass  $\mathcal{O}(10)$  MeV). At this scope, we modify the effective-interaction couplings by introducing an explicit dependence on the mediator mass  $m_a$  arising from the propagator. Consequently, the differential CEvNS cross sections for the NGI interactions read [66]

$$\left. \frac{d\sigma_{\nu\mathcal{N}}}{dT_{\mathcal{N}}} \right|^S (E_{\nu}, T_{\mathcal{N}}) = \frac{m_{\mathcal{N}} C_S^4}{4\pi(m_S^2 + 2m_{\mathcal{N}}T_{\mathcal{N}})^2} F_W^2(|\mathbf{q}|^2) \frac{m_{\mathcal{N}} T_{\mathcal{N}}}{E_{\nu}^2}, \quad (5)$$

$$\left. \frac{d\sigma_{\nu\mathcal{N}}}{dT_{\mathcal{N}}} \right|^V (E_{\nu}, T_{\mathcal{N}}) = \left[ 1 + \kappa \frac{C_V}{\sqrt{2}G_F Q_V^{\text{SM}} (m_V^2 + 2m_{\mathcal{N}}T_{\mathcal{N}})} \right]^2 \left. \frac{d\sigma_{\nu\mathcal{N}}}{dT_{\mathcal{N}}} \right|^{\text{SM}} \quad (6)$$

$$\left. \frac{d\sigma_{\nu\mathcal{N}}}{dT_{\mathcal{N}}} \right|^A (E_{\nu}, T_{\mathcal{N}}) = \frac{2m_{\mathcal{N}}}{2J+1} \frac{g_A^4}{(m_A^2 + 2m_{\mathcal{N}}T_{\mathcal{N}})^2} \left( 2 + \frac{m_{\mathcal{N}}T_{\mathcal{N}}}{E_{\nu}^2} - \frac{2T_{\mathcal{N}}}{E_{\nu}} \right) \tilde{S}^{\mathcal{T}}(|\mathbf{q}|^2), \quad (7)$$

$$\begin{aligned} \left. \frac{d\sigma_{\nu\mathcal{N}}}{dT_{\mathcal{N}}} \right|^T (E_{\nu}, T_{\mathcal{N}}) &= \frac{m_{\mathcal{N}}}{2J+1} \frac{g_T^4}{(m_T^2 + 2m_{\mathcal{N}}T_{\mathcal{N}})^2} \\ &\times \left[ \left( 2 - \frac{m_{\mathcal{N}}T_{\mathcal{N}}}{E_{\nu}^2} - \frac{2T_{\mathcal{N}}}{E_{\nu}} \right) \tilde{S}^{\mathcal{T}}(|\mathbf{q}|^2) + \left( 1 - \frac{T_{\mathcal{N}}}{E_{\nu}} \right) \tilde{S}^{\mathcal{L}}(|\mathbf{q}|^2) \right]. \quad (8) \end{aligned}$$

Note that the cross section for the vector interactions depends on the specific SM extension under consideration:  $\kappa = 1$  in the universal scenario, while  $\kappa = -1/3$  in the B – L model [67, 68]. The axial-vector and tensor cross section are written directly in terms of the fundamental coupling  $g_a$ , while the couplings  $C_a$  in the scalar and vector cross sections can be related to  $g_a$  at the quark level following the procedure of DM detection [69, 70] and are given by

$$C_S^2 \equiv g_S^2 \left( Z \sum_{q=u,d} \frac{m_p}{m_q} f_{T_q}^{(p)} + N \sum_{q=u,d} \frac{m_n}{m_q} f_{T_q}^{(n)} \right), \quad (9)$$

$$C_V^2 \equiv 3Ag_V^2. \quad (10)$$

Throughout this work, we assume that the new mediator  $a$  couples with equal strength to neutrinos, quarks and charged leptons. Under this assumption,  $g_a$  is defined as  $g_a = \sqrt{g_{\nu a} g_{qa}} = \sqrt{g_{\nu a} g_{\ell a}}$ , where  $g_{\nu a}$ ,  $g_{qa}$  and  $g_{\ell a}$  are the couplings between the mediator and neutrinos, quarks and leptons, respectively. In the previous expressions,  $m_p$  and  $m_n$  denote the proton and neutron masses, respectively, and  $m_q$  are the quark  $q$  masses, while  $f_{T_q}^{(p)}$  and  $f_{T_q}^{(n)}$  represent the quark mass contributions to the nucleon (proton and neutron) mass. Note that the expressions for the axial-vector and tensor mediated cross sections are spin dependent. For the latter two we have explicitly extracted their dependence on the total angular momentum,  $J$ , of the nucleus in the ground state. In the case of xenon nuclei, only the  $^{129}\text{Xe}$  and  $^{131}\text{Xe}$  isotopes have spin different from zero ( $J^{129}\text{Xe} = 1/2$  and  $J^{131}\text{Xe} = 3/2$ ), and therefore induce non-zero axial and tensor contributions. The respective abundances are 26.4% and 21.2%. The spin structure functions  $\tilde{S}^{\kappa}(|\mathbf{q}|^2)$ , where  $\kappa = \mathcal{L}, \mathcal{T}$ , account for longitudinal and transverse multipoles calculated using the Shell Model, and have been obtained following [71] as explained in Appendix B of [66]<sup>2</sup>. The pseudoscalar interaction is not considered in the following as it turns out to be negligible [71] for two reasons: first, its cross section is nuclear-spin suppressed and secondly, it is proportional to  $\frac{T_{\mathcal{N}}^2}{2E_{\nu}^2}$ , and hence also kinematically suppressed.

<sup>2</sup> Note that for the axial-vector interaction the longitudinal contribution is negligible in the CEvNS case [71], contrary to the upscattering scenario studied in [66].

### III. STATISTICAL ANALYSIS

We now proceed to discuss the implementation of the statistical analysis. We analyze the experimental data presented in Refs. [41, 42]. Both experiments utilize a dual-phase time-projection-chamber (TPC), and produce both scintillation photons (S1 signal) and ionization electrons (S2 signal). In the case of PandaX-4T two data sets were analyzed by the Collaboration, one corresponding to a paired S1 and S2 signal and one using only S2 (referred to as US2). In this paper, we use only the second data set (US2), since not enough information is provided by the experimental Collaboration for an accurate reproduction of the paired data.

The differential event rate is obtained by a convolution of the neutrino flux with the CE $\nu$ NS cross section

$$\frac{dR^{\text{X,P}}}{dT_{\mathcal{N}}} = \mathcal{A}^{\text{X,P}}(T_{\mathcal{N}}) \int dE_{\nu} \frac{d\phi}{dE_{\nu}} \frac{d\sigma_{\nu\mathcal{N}}}{dT_{\mathcal{N}}}, \quad (11)$$

where  $\frac{d\sigma_{\nu\mathcal{N}}}{dT_{\mathcal{N}}}$  refers to any of the expressions in Eq. (1) or Eqs. (5)–(8), and  $\mathcal{A}^{\text{X,P}}(T_{\mathcal{N}})$  is the experiment-dependent efficiency (X stands for XENONnT and P for PandaX-4T) which has been extracted from Refs. [41, 72] and Ref. [42] for XENONnT and PandaX-4T, respectively. The flux of  $^8\text{B}$  solar neutrinos,  $\frac{d\phi}{dE_{\nu}}$ , is taken from Refs. [73, 74] with the normalization defined in Ref. [45] ( $5.46 \times 10^6 \text{ cm}^{-2} \text{ s}^{-1}$ ). The data in Refs. [41, 42] is presented in bins of S2 (number of electrons,  $N_{e^-}$ ) in the case of XENONnT (PandaX-4T). Therefore, the events per bin are given by

$$R_i^{\text{X,P}} = c_i \mathcal{E}^{\text{X,P}} \int_i \frac{dR^{\text{X,P}}}{dn^{\text{X,P}}} dn^{\text{X,P}}, \quad (12)$$

where the integral is performed over the size of bin  $i$ , while  $n^{\text{X}} = \text{S2}$  and  $n^{\text{P}} = N_{e^-}$  for the case of XENONnT and PandaX-4T, respectively. Following the Collaborations, for XENONnT we consider 3 bins in the range [120, 500] photoelectrons (PE), while for PandaX-4T we consider 8 bins in the range [4, 8]  $N_{e^-}$ <sup>3</sup>. Note that our simulations do not account for resolution effects since no information about them is provided in the experimental papers. However, even without smearing, we are able to reproduce reasonably well the predicted event rates. Even under the same conditions (e.g. same assumptions on flux normalizations), we need to include the correction factors  $c_i$  in Eq. (12) in order to match our predictions with the best fit spectra presented in the experimental papers, see the first panel of Fig. 2 in Ref. [41] for XENONnT and the upper panel of Fig. 5 in Ref. [42] for PandaX-4T. These factors can be seen as effective efficiencies, included because we are performing a simplified analysis compared to what is done by the Collaborations. Indeed, we only use information on S2, while the experimental analyses rely on many more variables that are fitted simultaneously in a correlated way. The inclusion of these factors has nonetheless little effect on the NGI analyses, while in the case of the SM analyses it helped to better reproduce the results from the experimental Collaborations. Going back to Eq. (12),  $\mathcal{E}^{\text{X,P}}$  is the exposure at each experiment, i.e., 3.51 t $\times$ y for XENONnT and 1.04 t $\times$ y for PandaX-4T (US2), while the differential event rates are expressed through a change of variables according to

$$\frac{dR^{\text{X,P}}}{dn^{\text{X,P}}} = \frac{dR^{\text{X,P}}}{dT_{\mathcal{N}}} \frac{dT_{\mathcal{N}}}{dn^{\text{X,P}}}. \quad (13)$$

<sup>3</sup> These values correspond to a nuclear recoil energy range of [0.97, 5.10] keV in the case of XENONnT and [0.66, 1.19] keV for PandaX-4T.

In the case of XENONnT, the translation between nuclear recoil energy and the S2 signal is carried out through

$$n^X = S2 = T_N Q_y^X(T_N) g_2, \quad (14)$$

where  $g_2 = 16.9$  PE/electron and the charge yield  $Q_y^X(T_N)$  is taken from Ref. [75]. For PandaX-4T we use instead

$$n^P = N_{e^-} = T_N Q_y^P(T_N), \quad (15)$$

with the charge yield  $Q_y^P(T_N)$  given in Ref. [42].

The overall predicted number of events in a given bin  $i$  is eventually given by

$$N_i^{X,P} = R_i^{X,P} + \sum_k B_i^k, \quad (16)$$

where the spectra of the background components  $B_i^k$  are taken from Refs. [42, 72]. Regarding the experimental data measured by the two Collaborations,  $D_k^{X,P}$ , XENONnT has observed  $\sum_k D_k^X = 37$  events, accounting for both ionization and scintillation signals. In the case of PandaX-4T, 3 (332) events are observed for the paired (US2) signals. Our predictions are hence compared with the data  $D_k^{X,P}$  using

$$\chi_{X,P}^2 = \min_{\alpha, \vec{\beta}} \left\{ 2 \left( \sum_k N_k^{X,P} - D_k^{X,P} + D_k^{X,P} \ln \left( D_k^{X,P} / N_k^{X,P} \right) \right) + (\alpha / \sigma_\alpha)^2 + \sum_i (\beta_i / \sigma_{\beta_i})^2 \right\}, \quad (17)$$

where  $\alpha$  is a nuisance parameter with  $\sigma_\alpha = 12\%$  uncertainty accounting for the  $^8\text{B}$  flux prediction, and  $\vec{\beta}$  and  $\sigma_{\beta_i}$  are the remaining nuisance parameters and uncertainties of the experiments. All nuisance parameters are included as normalization factors on the components in Eq. (16). In the case of XENONnT we include an uncertainty of 5% on our signal prediction related to the fiducial volume. In addition, the background components receive the following uncertainties: a 4.8% uncertainty for accidental coincidence (AC), a 50% uncertainty for the neutron-related background and a 100% uncertainty for the electron recoil (ER) background, both being subleading compared to AC. For PandaX-4T we use a 22% uncertainty on the signal prediction due to data selection and interaction modelling. We further include uncertainties of 31% and 23% for the cathode and micro-discharges (MD) background components. Finally, let us note that we also perform a combined analysis of XENONnT and PandaX-4T data. In this case, the correlated uncertainty on the neutrino flux is included only once.

In Fig. 1 we demonstrate the distributions of signal and background events as a function of the number of ionized electrons, for PandaX-4T (left panel) and XENONnT (right panel). In the case of PandaX-4T, the green histogram represents the radioactivity on the cathode electrode (CE) while the micro-discharging background is summed over the cathode background in yellow. The magenta histogram finally accounts for the CE $\nu$ NS prediction plus both the CE and MD backgrounds, as given by the experimental Collaboration. Our total prediction is given as a blue line, and has to be compared to the magenta histogram. Experimental data are also shown together with their error bars. In the case of XENONnT, the CE $\nu$ NS signal is represented in light green on top of the backgrounds, indicated by light purple (AC) and light blue (electron recoil). The neutron recoil background is barely visible in the plot, but it is also included.

#### IV. RESULTS

In this section we present the results of our analyses first concerning SM physics ( $^8\text{B}$  solar neutrino flux and a determination of the weak mixing angle) in Sec. IV A and then for new light mediators in Sec. IV B.



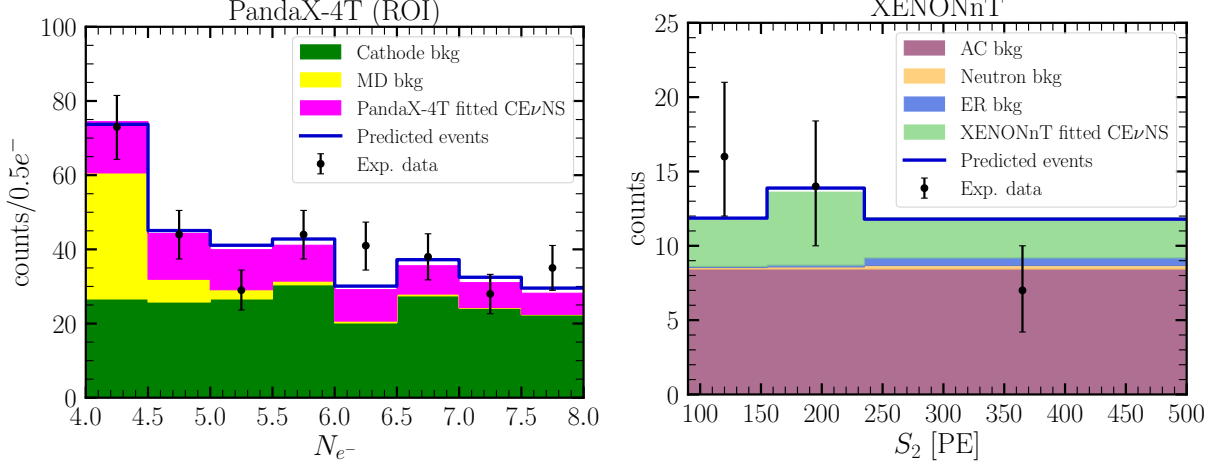


FIG. 1: Distribution of signal and background fitted events for PandaX-4T (left panel) and XENONnT (right panel). The colored histograms are given in the experimental papers, while the blue lines are our predictions of CE $\nu$ NS plus background events for the two analyses. The measured events together with the error bars are also shown for each experiment.

### A. SM physics: weak mixing angle and $^8\text{B}$ flux

The experimental results announced by the XENONnT and PandaX-4T Collaborations allow for a measurement of the  $^8\text{B}$  solar neutrino flux through its CE $\nu$ NS-induced signal. Both Collaborations claim agreement with the standard solar model prediction and with other dedicated solar neutrino experiments, indicating a constraint of  $\Phi_\nu^{^8\text{B}} = (4.7^{+3.6}_{-2.3}) \times 10^6 \text{ cm}^{-2} \text{ s}^{-1}$  at 68% confidence level (CL) in the case of XENONnT and  $\Phi_\nu^{^8\text{B}} = (8.4 \pm 3.1) \times 10^6 \text{ cm}^{-2} \text{ s}^{-1}$  at 68% CL for PandaX-4T, obtained using a combined analysis of paired and US2 data. To test our statistical analysis, we also extract the constraints on the  $^8\text{B}$  solar neutrino flux for both experiments separately, and from a combined analysis. At this scope, we assume the flux-weighted CE $\nu$ NS cross section as predicted in the SM, fixing for this analysis  $\sin^2 \theta_W = 0.23857$ . The reduced  $\chi^2$ -profiles are shown in the left panel of Fig. 2: the green dashed curve corresponds to PandaX-4T, the blue dot-dashed one to XENONnT, while the magenta plain one depicts the combined result. At  $1\sigma$  CL our results read

$$\Phi_\nu^{^8\text{B}} = (7.7^{+7.0}_{-5.9}) \times 10^6 \text{ cm}^{-2} \text{ s}^{-1} \quad (\text{PandaX-4T}), \quad (18)$$

$$\Phi_\nu^{^8\text{B}} = (5.7^{+3.2}_{-2.8}) \times 10^6 \text{ cm}^{-2} \text{ s}^{-1} \quad (\text{XENONnT}), \quad (19)$$

$$\Phi_\nu^{^8\text{B}} = (6.1^{+2.8}_{-2.7}) \times 10^6 \text{ cm}^{-2} \text{ s}^{-1} \quad (\text{combined}). \quad (20)$$

Our result does not agree very well with the official result from PandaX-4T stated above. However, since we use a reduced data set (only US2) a weaker result on the solar neutrino flux could be expected. We verified, in any case, that for our best fit value the overall number of events (we obtain 69) lies within the stated  $1\sigma$  interval of the Collaboration for the US2-only analysis, which is [42]:  $92 \pm 34$ .

As anticipated in Sec. II, one relevant SM parameter entering the CE $\nu$ NS cross section is the weak mixing angle,  $\sin^2 \theta_W$ . The observation of CE $\nu$ NS data at DM direct detection experiments allow to determine this parameter at low energy, i.e., at a renormalization scale  $\mu \simeq \mathcal{O}(10)$  MeV, corresponding to the typical momentum transfer exchanged in the process. A variation in the value of  $\sin^2 \theta_W$  would affect the overall normalization of the CE $\nu$ NS event rates. We perform a statistical analysis, this time fixing  $\Phi_\nu^{^8\text{B}} = 5.46 \times 10^6 \text{ cm}^{-2} \text{ s}^{-1}$  (but adding the associated 12% uncertainty),

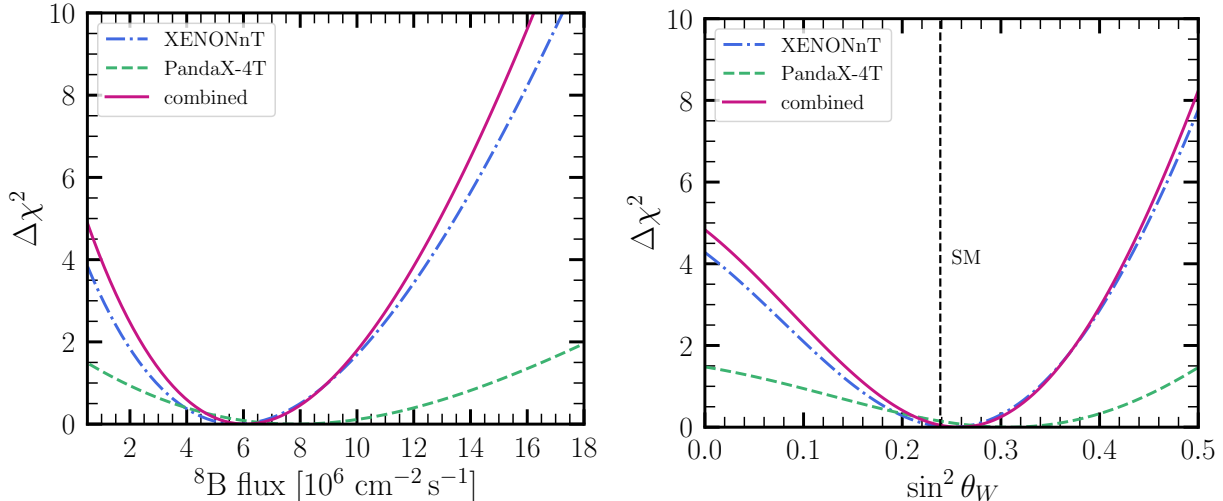


FIG. 2: Reduced  $\chi^2$ -profiles for the determination of the  $^8\text{B}$  solar neutrino flux (left) and of the weak mixing angle (right) for the PandaX-4T (green dashed), XENONnT (blue dot-dashed) and the combined (magenta plain) analyses. In the right plot, the vertical black dashed line indicates the SM value from the RGE running in the  $\overline{\text{MS}}$  renormalization scheme.

and we extract the following best fit values and  $1\sigma$  uncertainties

$$\sin^2 \theta_W = 0.30^{+0.16}_{-0.21} \quad (\text{PandaX-4T}), \quad (21)$$

$$\sin^2 \theta_W = 0.25^{+0.09}_{-0.10} \quad (\text{XENONnT}), \quad (22)$$

$$\sin^2 \theta_W = 0.26^{+0.08}_{-0.09} \quad (\text{combined}). \quad (23)$$

Figure 2 (right) shows the reduced  $\chi^2$ -profiles for the determination of the weak mixing angle, for the two separate data sets and for the combined analysis. Moreover, we show in Fig. 3 the best fit values together with the  $1\sigma$  error bars, as a function of the renormalization scale. For comparison, the plot additionally shows the RGE evolution in the SM (coral dashed line), calculated in the  $\overline{\text{MS}}$  renormalization scheme [76] as well as other existing constraints at different energy scales [77–81]. Let us note that a determination of  $\sin^2 \theta_W$  from recent direct detection data was already performed in [53] leading to similar results in the case of XENONnT data. However, in the case of PandaX-4T our result differs from Ref. [53]. This might be due to the fact that our Eq. (13) does not agree with Eq. (7) in Ref. [53], since the charge yield  $Q_y$  also depends on the nuclear recoil energy and hence  $dT_N/dn$  is not  $1/Q_y$  as assumed in this reference.

Complementarity with other CEvNS measurements is particularly evident, for instance those from COHERENT CsI and liquid argon data [82], from the Dresden-II reactor experiment [83, 84], represented in gray, and from a combination of different electroweak measurements [85]. Even though not shown in the plot to avoid overcrowding the figure, other low-energy measurements have been extracted from spallation source and reactor CEvNS data [63, 86, 87], also in combination with data from atomic parity violation and parity-violating electron scattering on different nuclei [88, 89]. Sensitivities from elastic scattering off atomic electrons at IsoDAR [90], neutrino-electron scattering events at SBND [91] and DUNE [92] further complete the picture of low-energy  $\sin^2 \theta_W$  determinations. DM direct detection facilities, despite their different primary scientific goal, can still provide valuable information on the value of the neutrino neutral-current interaction at low energy. While the current error bars of the measurements at DM direct detection experiments are still large compared to other determinations at higher energy scales, these novel measurements lie



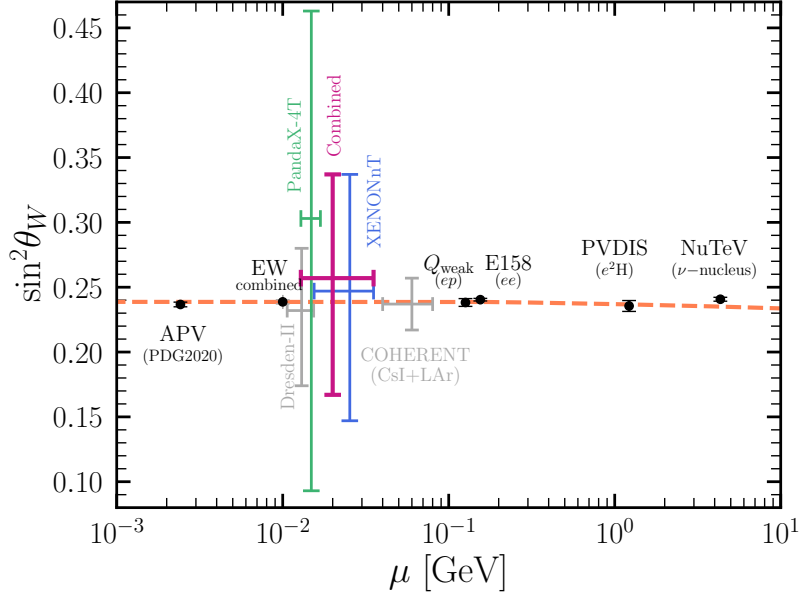


FIG. 3: Weak mixing angle running in the SM for the  $\overline{\text{MS}}$  renormalization scheme (coral dashed line) as a function of the renormalization scale. Our  $1\sigma$  determinations are shown in green for PandaX-4T, blue for XENONnT and magenta for the combined analysis. Measurements from other experiments [64, 78–82, 84, 85] are also shown for comparison.

in a part of parameter space previously unexplored. Forthcoming data and improved statistics at DM facilities will allow to extract the value of  $\sin^2\theta_W$  at  $\mu \simeq \mathcal{O}(10)$  MeV with higher precision, in full complementarity with dedicated CE $\nu$ NS experiments.

## B. Neutrino generalized interactions

Next we discuss the constraints obtained for new neutrino interactions in the presence of different light mediators, as defined in Sec. II B. In these analyses, we have kept the value of the weak mixing angle fixed at its SM value,  $\sin^2\theta_W = 0.23857$  and the normalization of the  $^8\text{B}$  solar neutrino flux  $\Phi_\nu^{^8\text{B}} = 5.46 \times 10^6 \text{ cm}^{-2} \text{ s}^{-1}$ .

The results are shown in Fig. 4, where we show the contours at 90% CL for XENONnT (upper left panel), PandaX-4T (upper right panel), and for the combined analysis (lower panel). The magenta, gold, red and green lines correspond to the analyses of tensor, axial, vector B-L, and scalar interactions as introduced in Sec. II B. The light blue shaded region denotes the excluded region in the case of universal vector interactions. We chose this format to highlight the fact that a region in the form of a thin band remains allowed in this scenario in the case of the XENONnT and the combined analysis. This degeneracy appears due to a destructive interference between the SM and the new vector couplings in the weak nuclear charge. In the PandaX-4T analysis we expect this degeneracy to appear below the currently excluded region once more statistics becomes available. Moreover, this cancellation can not occur in the case of the B-L model, due to the particle charges under the  $U(1)_{\text{B-L}}$  symmetry. As expected, the bounds on the spin-dependent axial and tensor mediators are much weaker than those for scalar and vector mediators.

In order to put our results into context, in Fig. 5 we compare our bounds with other existing constraints on the same types of interactions. We show our constraints obtained from the combined analysis of XENONnT + PandaX-4T (in blue, at 90% CL) for the scalar, vector universal, vector

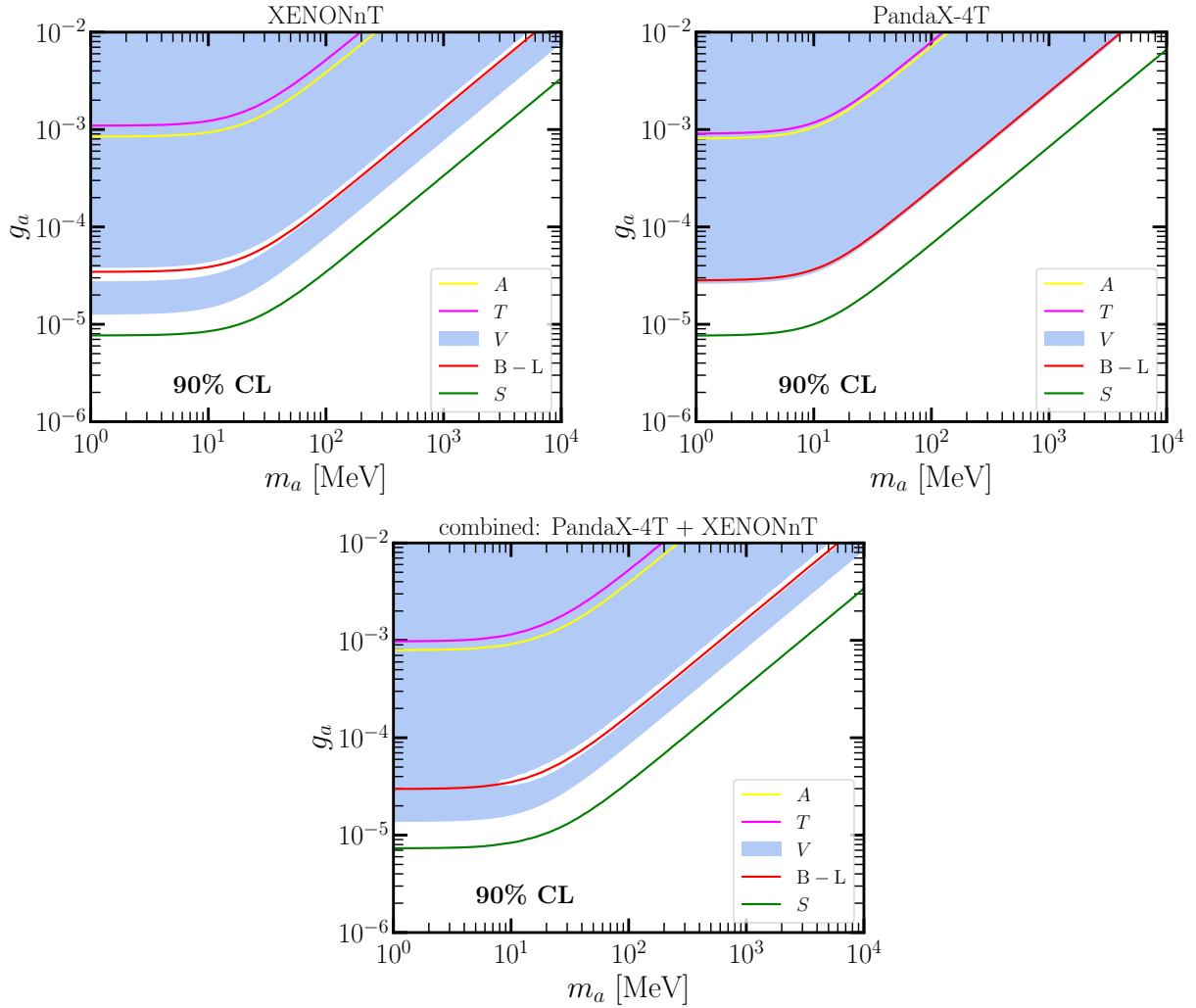


FIG. 4: The 90% CL exclusion limits for new neutrino generalized interactions with light mediators obtained from the analysis of XENONnT data (upper left), PandaX-4T data (upper right), and from the combined analysis (lower panel).

B-L and axial-vector interactions. Additionally, we show existing limits from other CEvNS data, in particular COHERENT [82, 93], CONUS [94, 95], and CONNIE [96, 97]; from elastic neutrino-electron scattering data at BOREXINO [98], CHARM-II [99] and TEXONO [99, 100]; from a combined analysis of PandaX-4T, XENONnT and LZ electron recoil data [36, 101]; from beam-dump and fixed-target experiments (including E141 [102], E137 [103], E774 [104], KEK [105], Orsay [106–108],  $\nu$ -CAL I [109–112], CHARM [113, 114], NOMAD [115], PS191 [116, 117], A1 [118] and APEX [119]); from colliders (BaBar [120, 121] and LHCb [122]); from NA64 [123–125]. To recast some of the bounds between the different interactions we used the DarkCast package [126, 127]. Note also that some of the bounds are a novel result in this work: we have recomputed the COHERENT axial-vector bound following [82], however including only the analysis of CsI, since the  $^{40}\text{Ar}$  nucleus is even. The combined XENONnT + PandaX-4T + LZ E $\nu$ ES bounds for the scalar, the universal vector and the axial-vector, not shown in [36], have also been computed specifically for this work. Similarly, the TEXONO bound on the scalar interaction is a new result. Finally, astrophysical and cosmological bounds are also particularly relevant for low-mass mediators. Even though strongly model-dependent and thus requiring a tailored analysis, we indicate with  $N_{\text{eff}}$  the

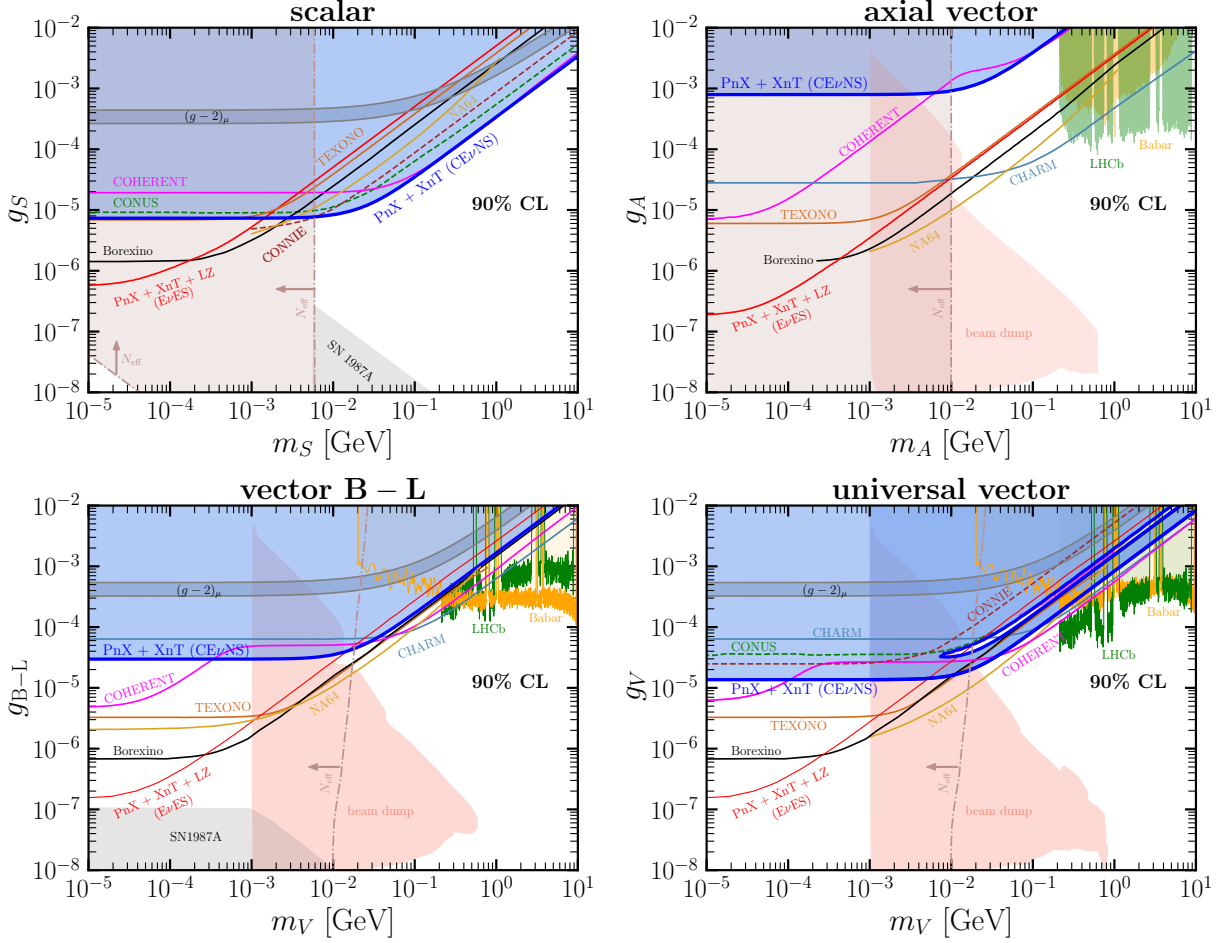


FIG. 5: The 90% CL excluded regions for some interactions obtained from the combined analysis of PandaX-4T and XENONnT data (blue): scalar (upper, left), axial-vector (upper, right), vector B-L (lower, left) and universal vector (lower, right). Existing bounds from other searches are also shown for comparison.

regions potentially in conflict with BBN [128–130] and CMB [131] and with SN1987A those in conflict supernova [132–136] data.

As can be seen, for the case of scalar-mediated processes, the constraints extracted in this work dominate for  $6 < m_S < 150$  MeV, improving previous results from dedicated CE $\nu$ NS experiments such as CONUS, CONNIE and COHERENT, while for  $m_S > 150$  MeV the sensitivity becomes similar to COHERENT. It is also noteworthy that the present sensitivities are the leading ones among the CE $\nu$ NS-based measurements in the region that is not in conflict with astrophysics. For the case of axial-vector interactions, as previously noted, the nuclear spin-suppression leads to poor sensitivities compared to e.g., the E $\nu$ ES-induced constraints from TEXONO, CHARM and Borexino as well as to those coming from a combined analysis of electron recoils at PandaX-4T, XENONnT and LZ. However, if the axial mediator coupling to quarks is assumed to be different from the coupling to leptons, then the present results —although very weak— are dominating in the mass range  $6 < m_A < 80$  MeV, while for larger masses they eventually become similar to COHERENT. Finally, focusing on the B-L and universal vector mediator models, the present sensitivities are not improving upon existing constraints, though being almost competitive in some regions of the parameter space. When compared to other CE $\nu$ NS constraints like COHERENT,

XENONnT and PandaX-4T offer a slight improvement in the range  $0.1 < m_V < 30$  MeV ( $0.3 < m_V < 40$  MeV) in the B-L (universal vector) case, in a region that is however in tension with astrophysical observations.

## V. CONCLUSIONS AND OUTLOOK

Indications of  $^8\text{B}$  solar neutrinos inducing  $\text{CE}\nu\text{NS}$  at DM direct detection experiments have recently been reported. Motivated by these results, we have analyzed the first  $\text{CE}\nu\text{NS}$  data collected by the PandaX-4T and XENONnT experiments. We have performed SM and new physics analyses showing that these data can be used to determine the weak mixing angle and to place a bound on the solar  $^8\text{B}$  neutrino flux. The obtained constraints on the weak mixing angle lie at a low energy scale, falling between the Dresden-II and COHERENT measurements. In addition, we have placed bounds on the mass and coupling of several light mediators, focusing on those with scalar, vector (universal and B-L), axial-vector, and tensor interactions. We have found that even with these first data we can place competitive bounds in some regions of parameter space, specially in the case of scalar and vector interactions, complementing other experimental probes including neutrino scattering data, beam dump and collider searches. In this paper we have focused on universal interactions, i.e., assuming equal couplings for all neutrino species, charged leptons and quarks. In these scenarios the bound is independent of the composition of the neutrino flux at the detector. In a future work we plan to extend the analysis including searches with non-universal and flavor-dependent couplings. More data is expected to be released in the future, from the DM experiments analyzed in this paper and also from the LZ experiment, which will allow us to further improve the bounds obtained here.

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